Sources for the EBL and cosmology studies

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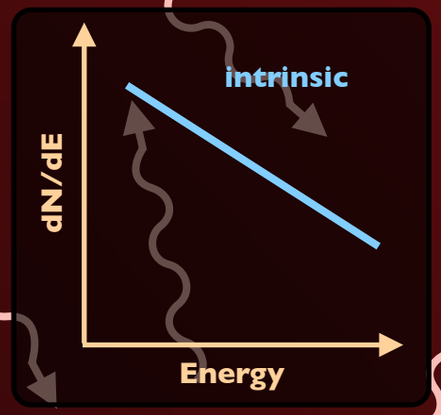
Contents

- Extragalactic Background Light
- Constraints so far
- Pile-ups at high energy: what are they?
- Steady sources for EBL/cosmology studies
- Hubble constant issue

AGN

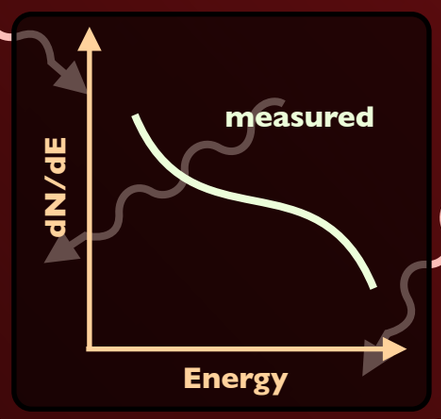


Stars and **Dust**
in Galaxies

HE/VHE γ-Rays

UV/O/IR
Photons

e+ e-



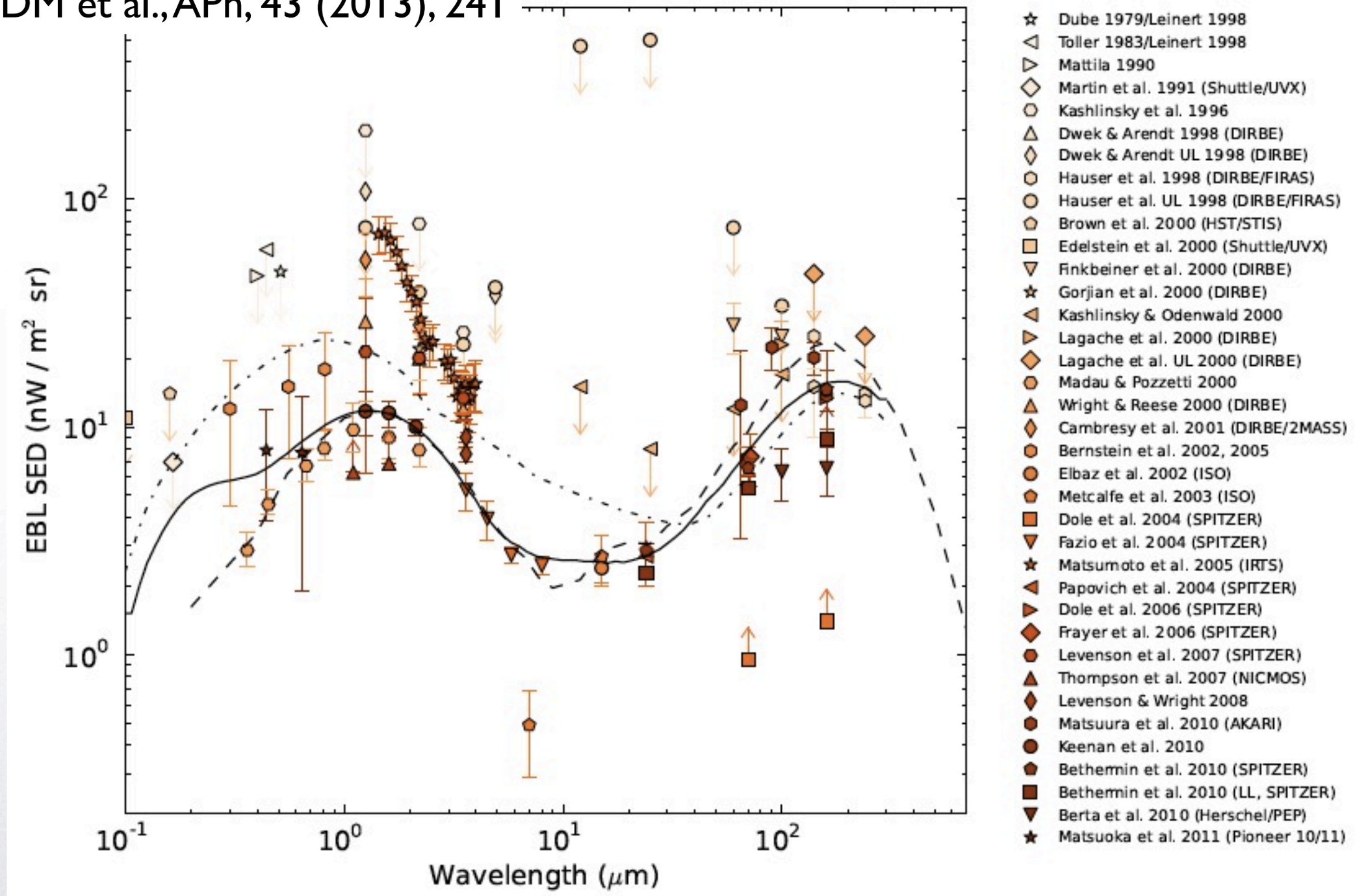
$$E_\gamma E_{\mathrm{EBL}} \approx 4(m_e c^2)^2 \approx 1\,\mathrm{MeV}^2$$
$$E_{\mathrm{EBL}} \sim \mathrm{eV} \rightarrow E_\gamma \sim \mathrm{TeV}$$

Nikishov (1962), Jelley (1966), Gould & Schreder (1966)

slide from M Raue

The EBL energy density

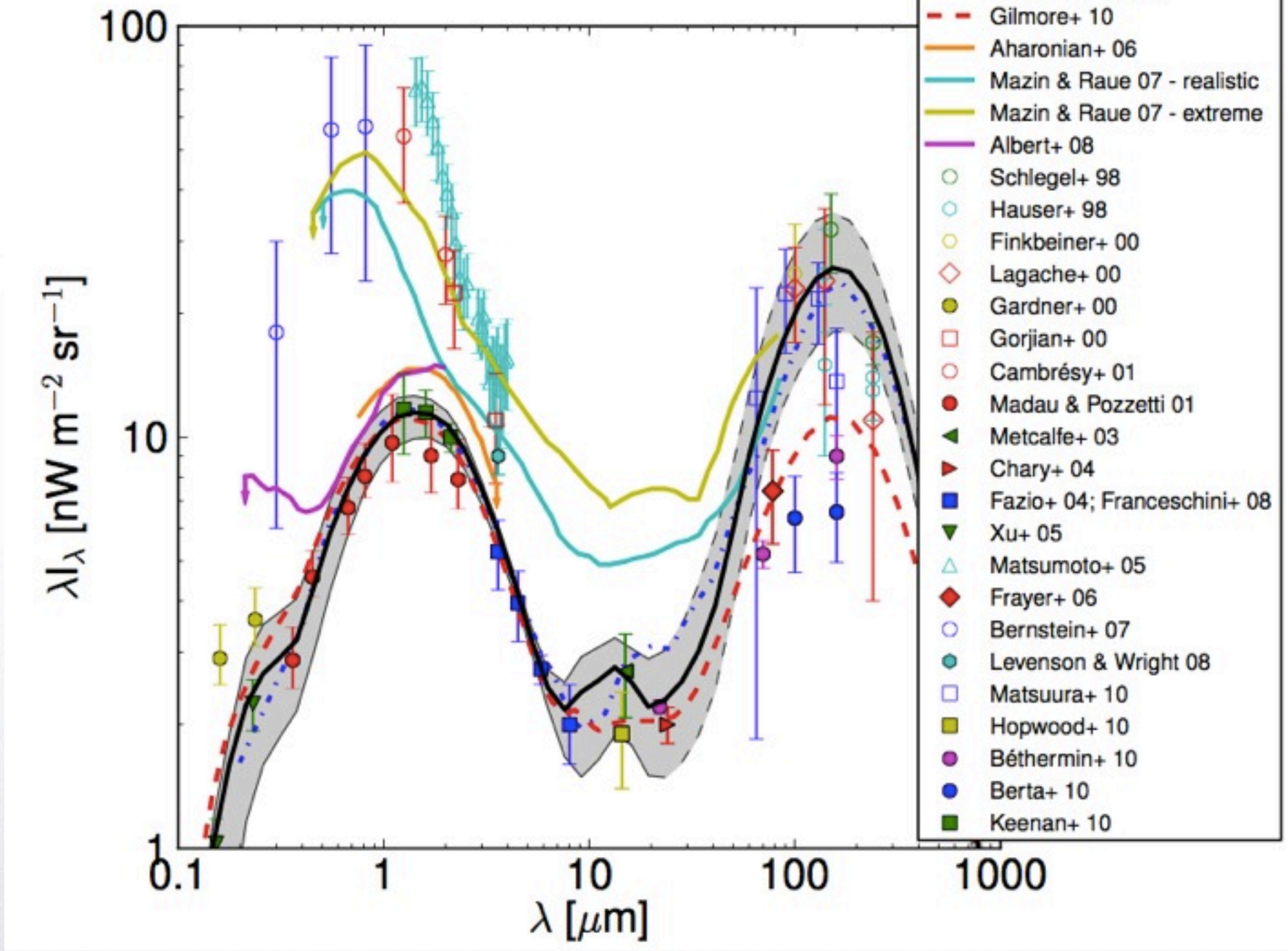
DM et al., APh, 43 (2013), 241



Sources for the EBL studies

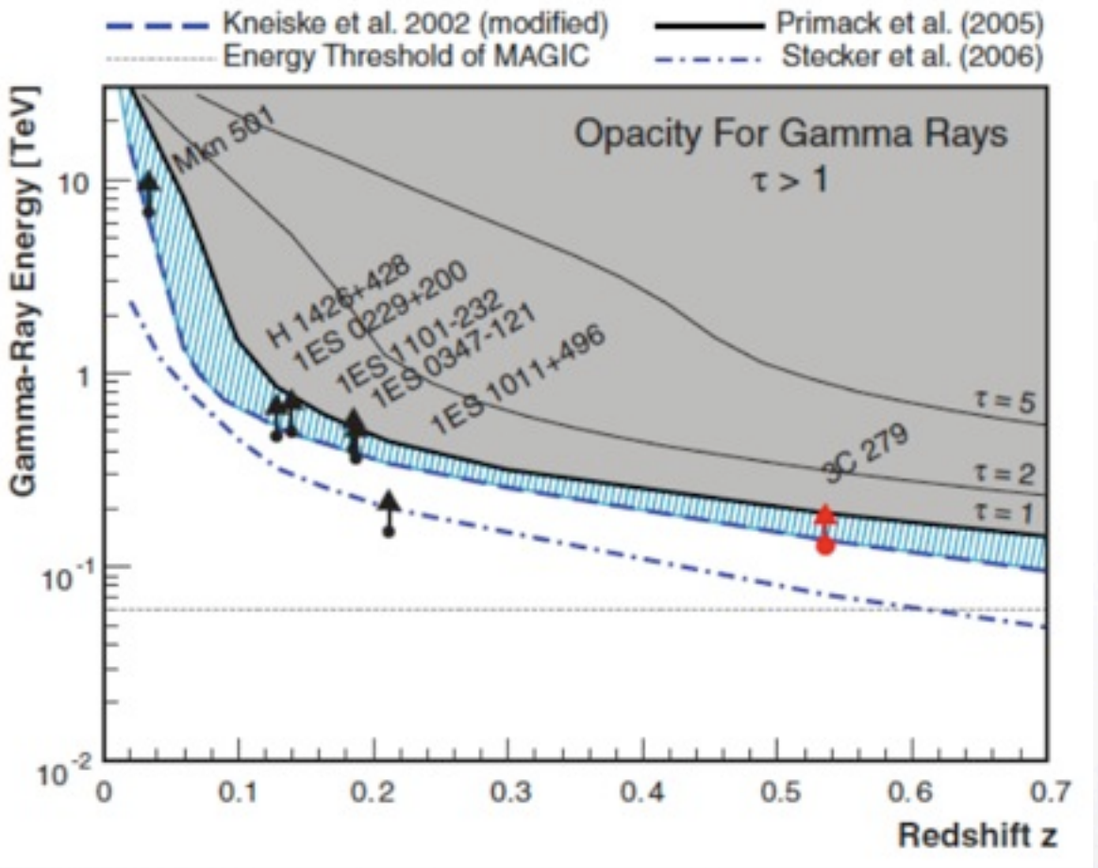
The EBL energy density

Dominguez et al., MNRAS 410, 2556 (2010)

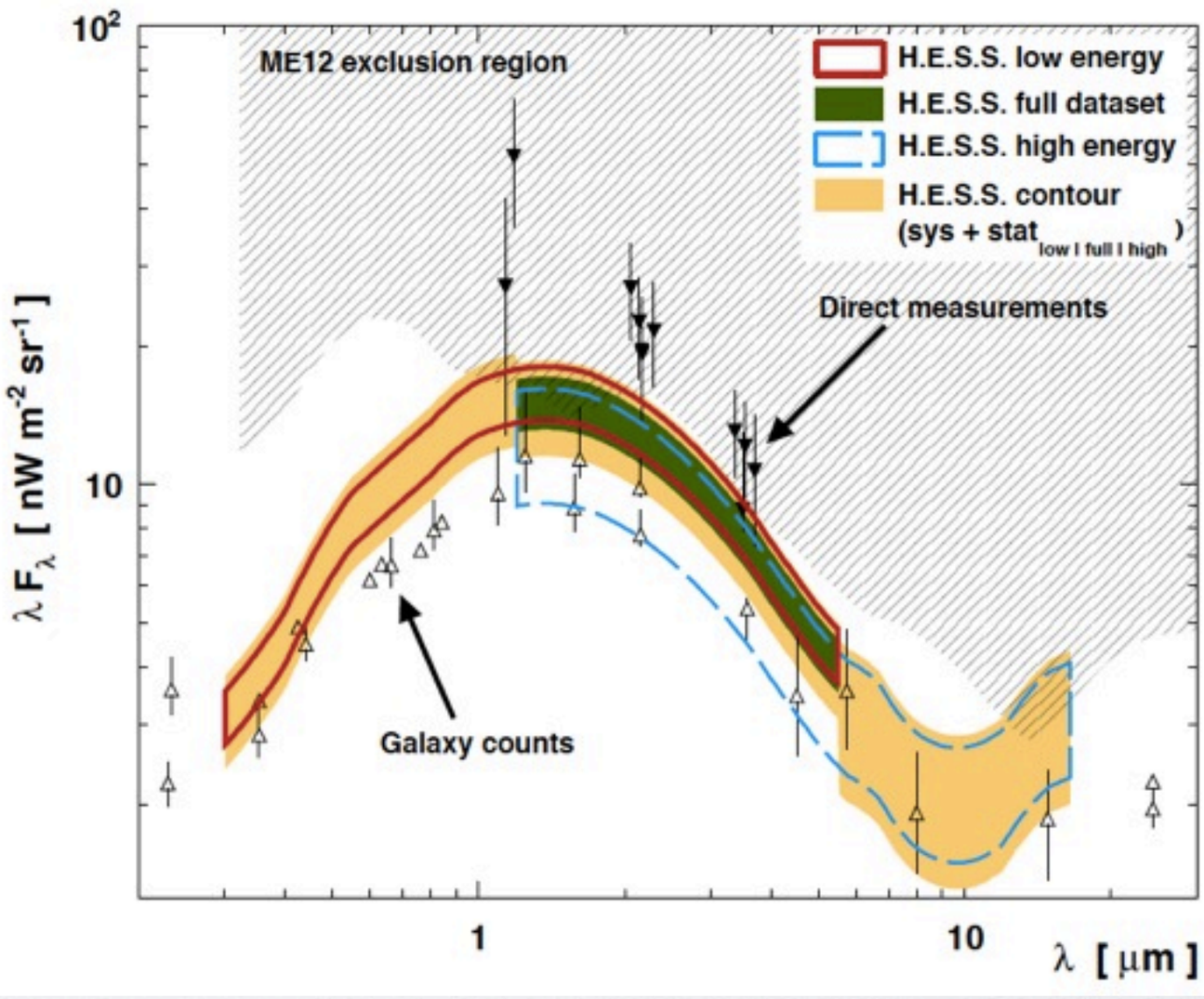


The EBL energy density

MAGIC, Science 320,1752 (2008)



H.E.S.S., A&A 550, 11 (2013)

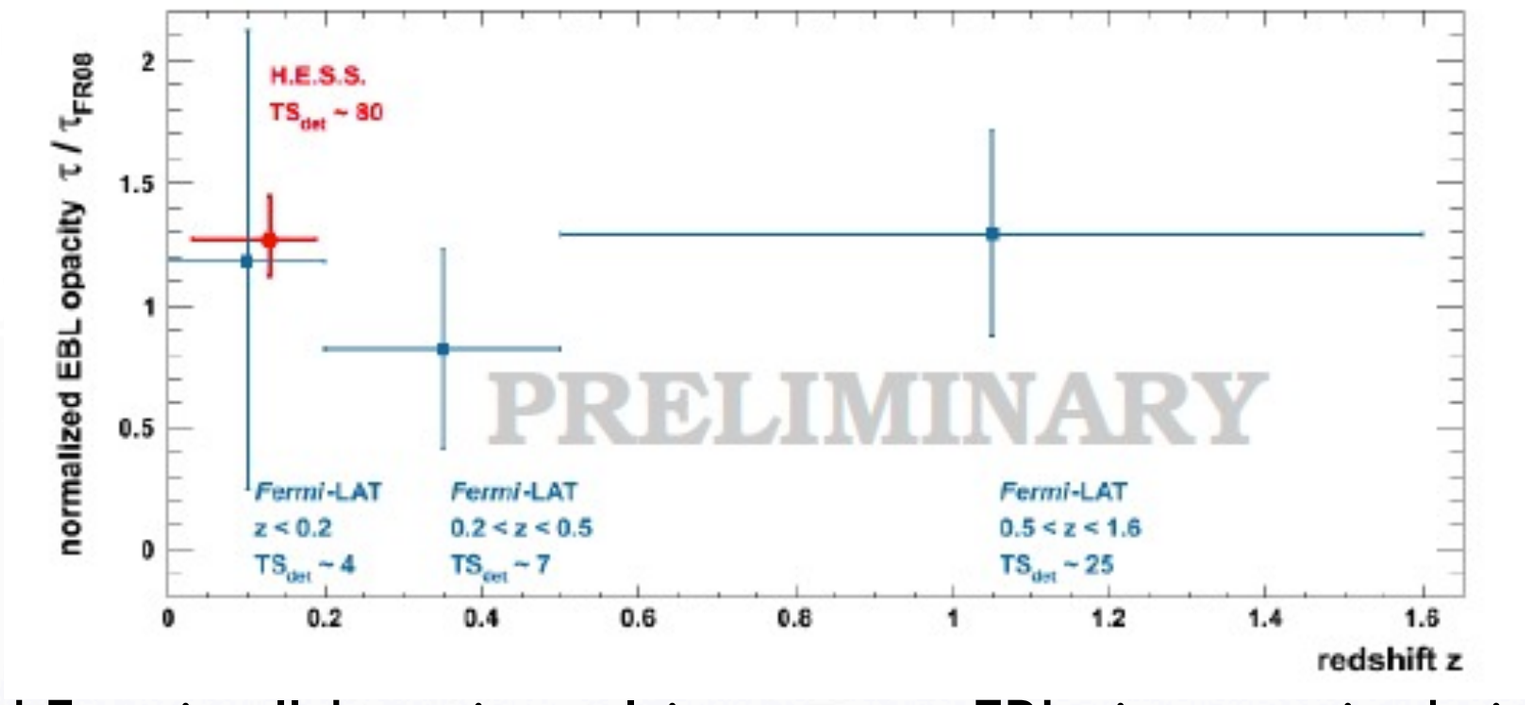


Not much more EBL than the one from the resolved galaxies

Combined Fermi-LAT and H.E.S.S. limits

Fermi/LAT, Science (2012) 338, 1190

HESS, A&A (2013) 550, 11

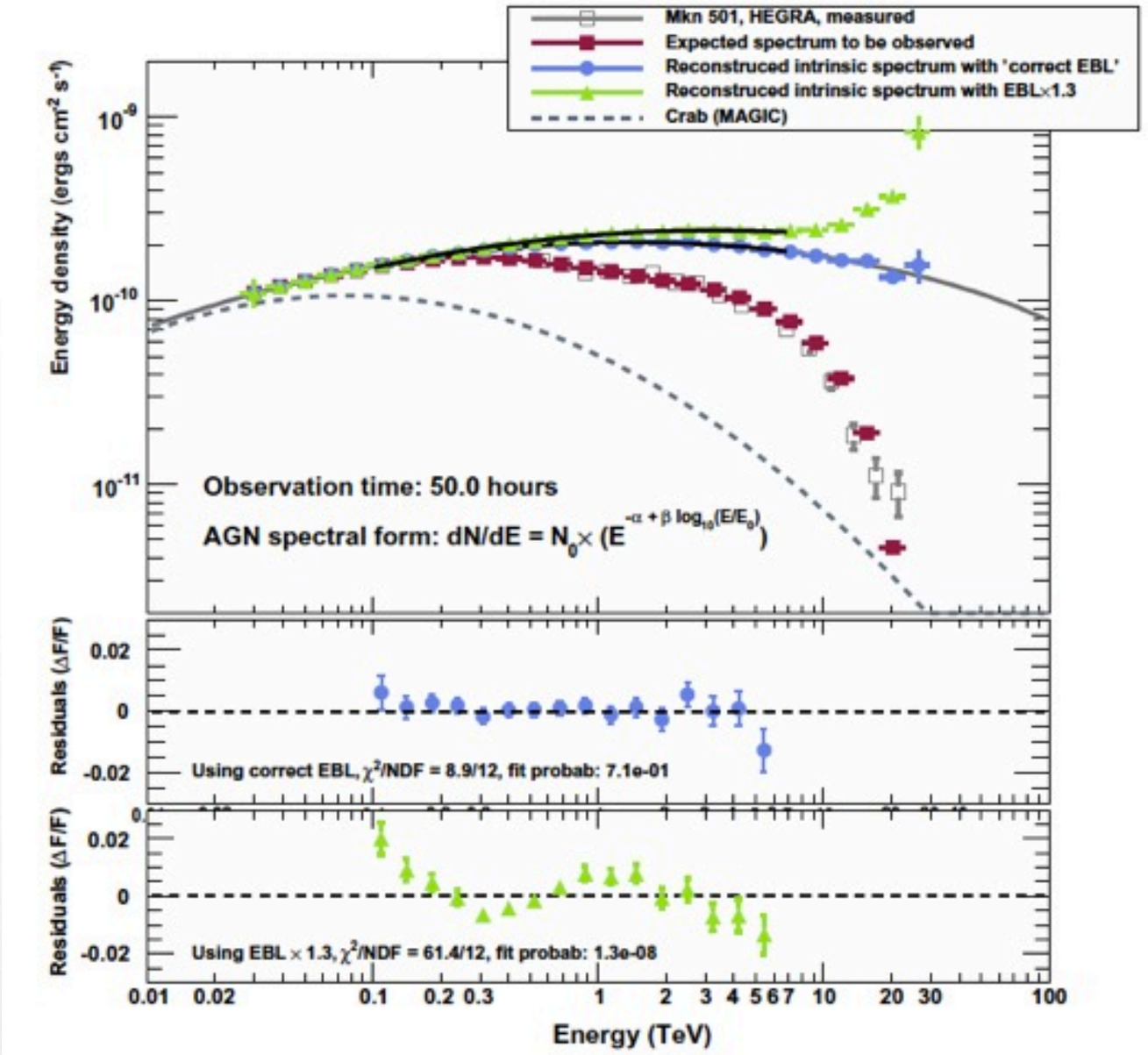


- HESS and Fermi collaborations claim to see a EBL signature in their blazar spectra.
- Assumption that there is no EBL can be excluded with 8σ and 6σ by HESS and Fermi, respectively.
- The EBL model which best fits the data (i.e. make the spectra smooth) is very close to recent established EBL models

Wiggles in spectra of AGN in very high state

precision EBL at Mid infrared

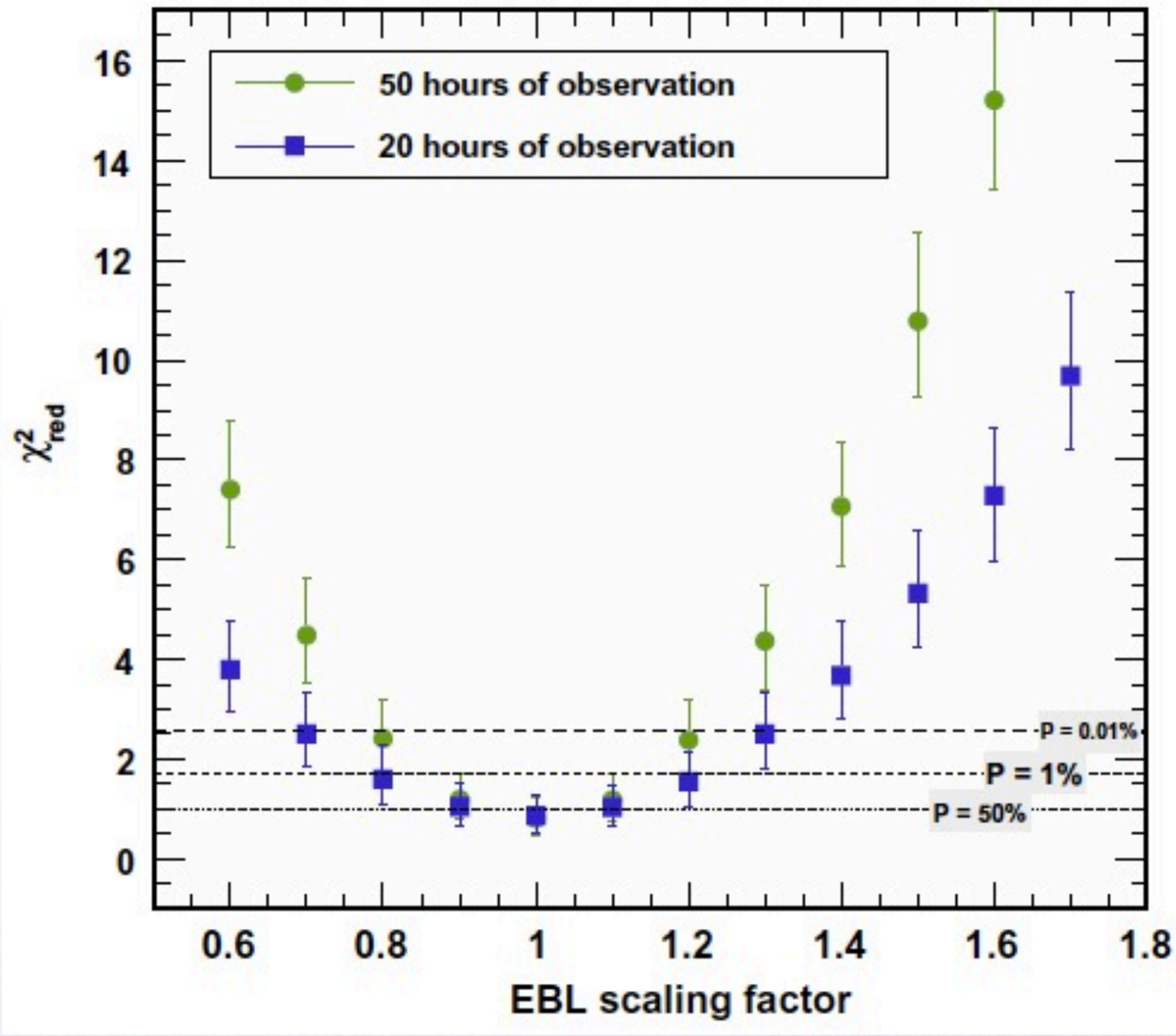
- Idea: if the measurement is precise enough, the spectral form will be sensitive to "wrong" EBL assumptions, and "wiggles" in the reconstructed spectrum would appear
- Here an example of Mrk 501 simulation:
  - **blue points** obtained with proper EBL model, no wiggles in residuals
  - **green points** obtained with a wrong (by 30%) EBL, wiggles clearly seen in the residuals



DM et al., APh 43 (2013), 241

Precision of EBL determination

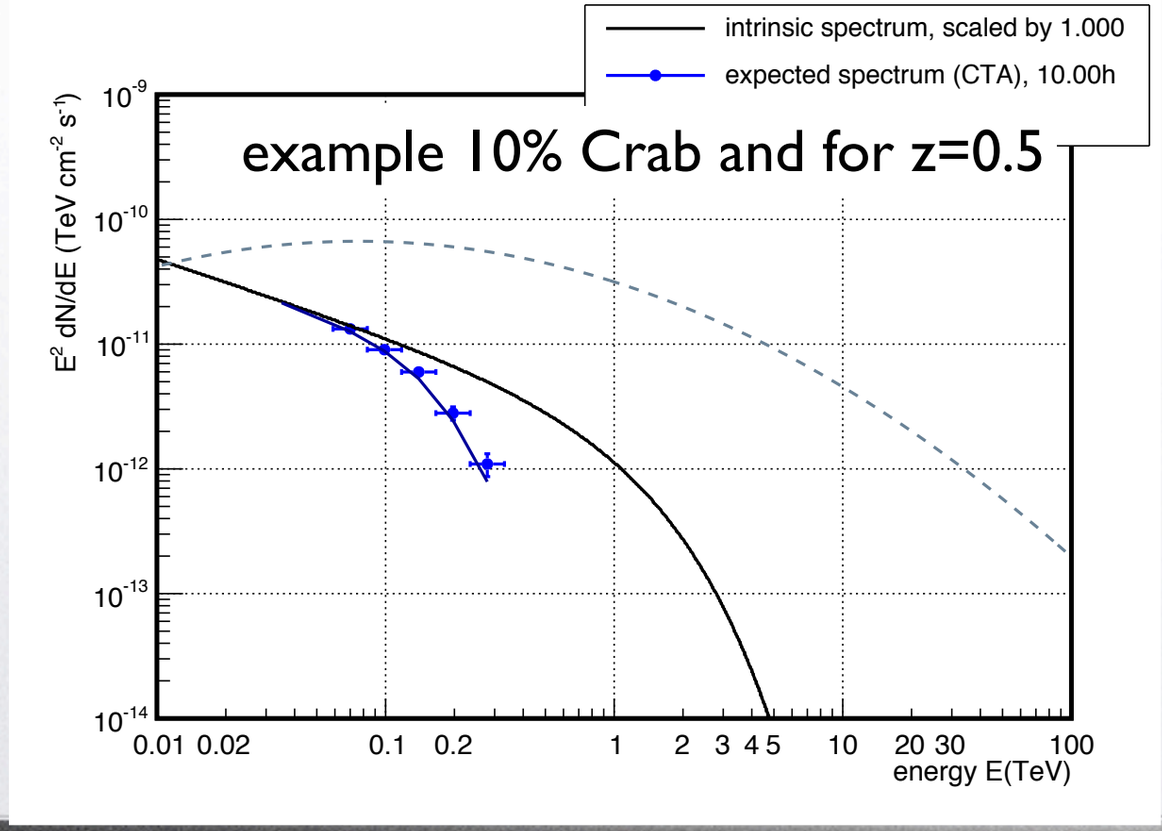
- Precision of the method is between 15 and 30%, depending on the assumed exposure of the AGN in flaring state

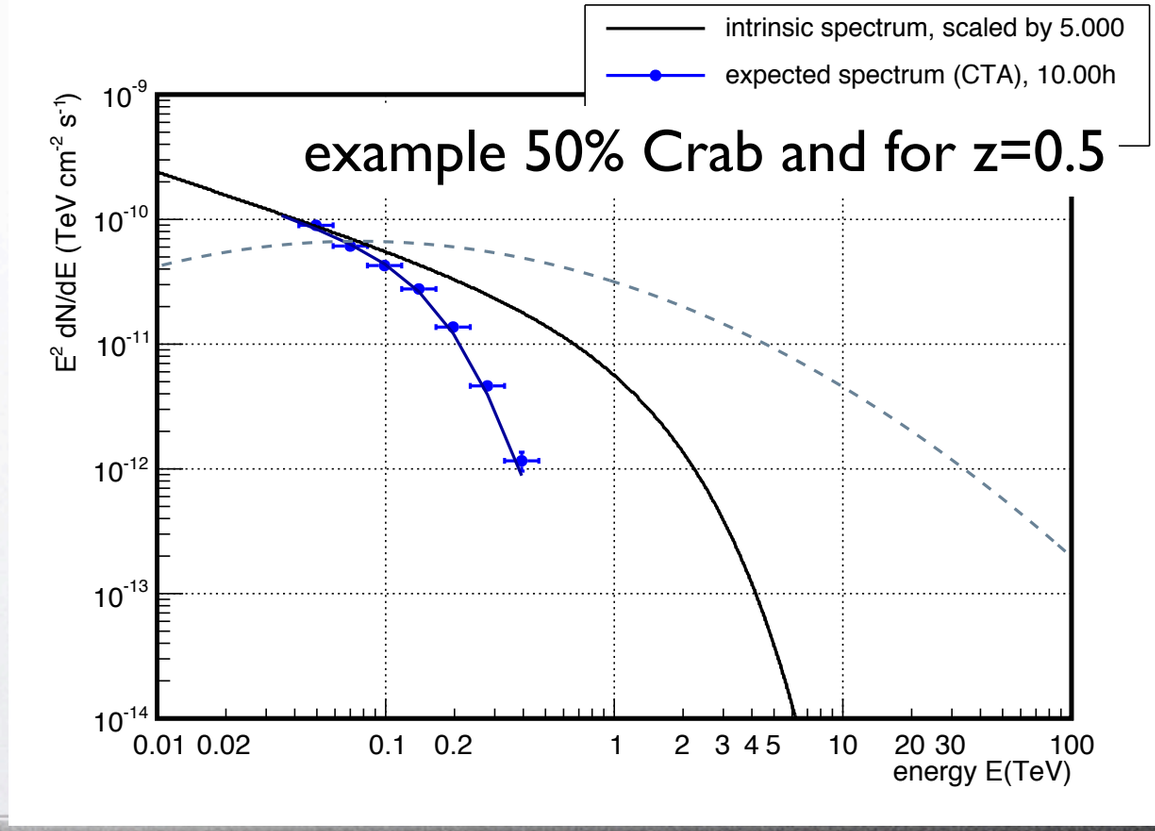


DM et al., APh 43 (2013), 241

Precision of EBL determination

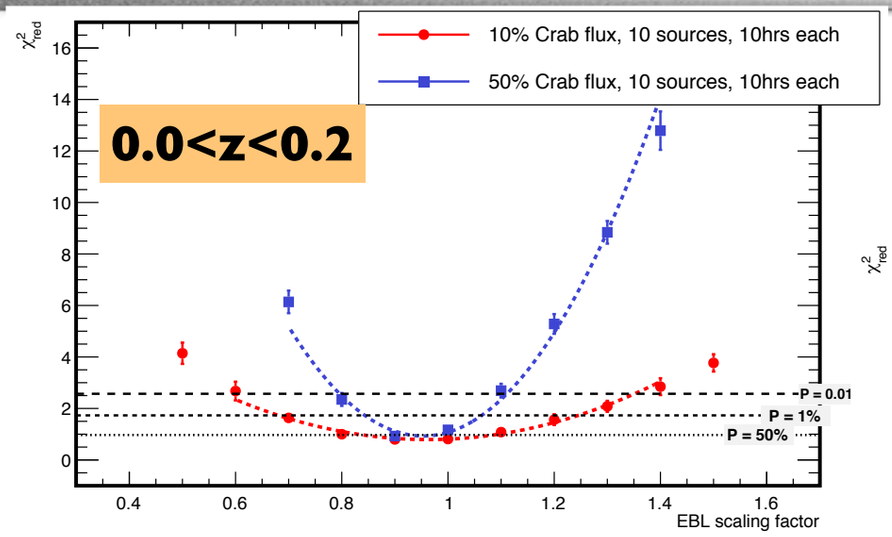
- Assume some 50 AGN at different redshifts, 0 < z < 1, are flaring for 10 hrs each; assume their spectra are known (through fit to intrinsic part of the spectrum where there is no EBL effect)
- Simulated two cases:
  - average flux during flaring episodes is 10% Crab, average index -2.6, intr. cutoff at 1 TeV and
  - average flux during flaring episodes is 50% Crab
- Fit resulting de-absorbed spectrum using $\chi 2$ fit

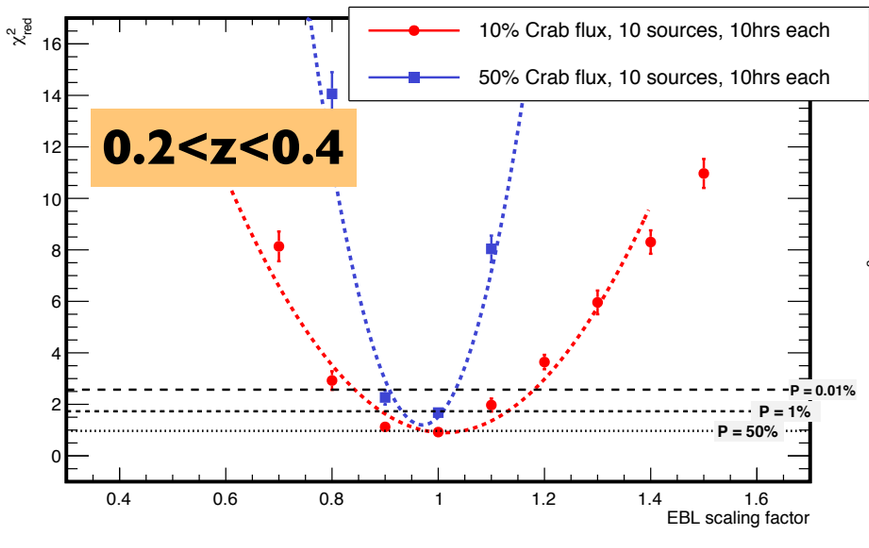


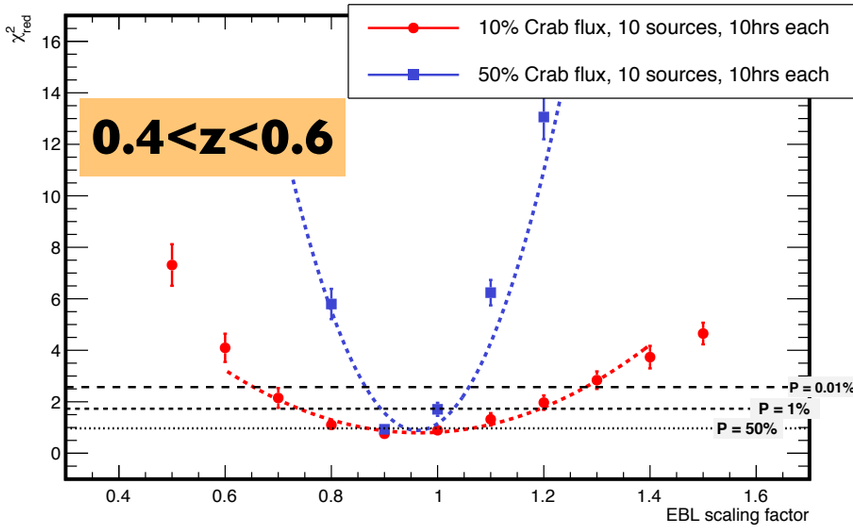


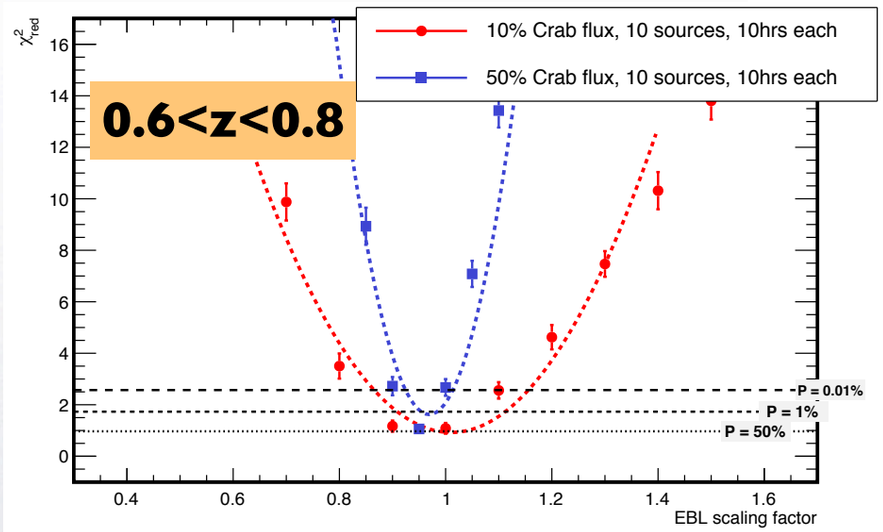
Precision of EBL determination (update)

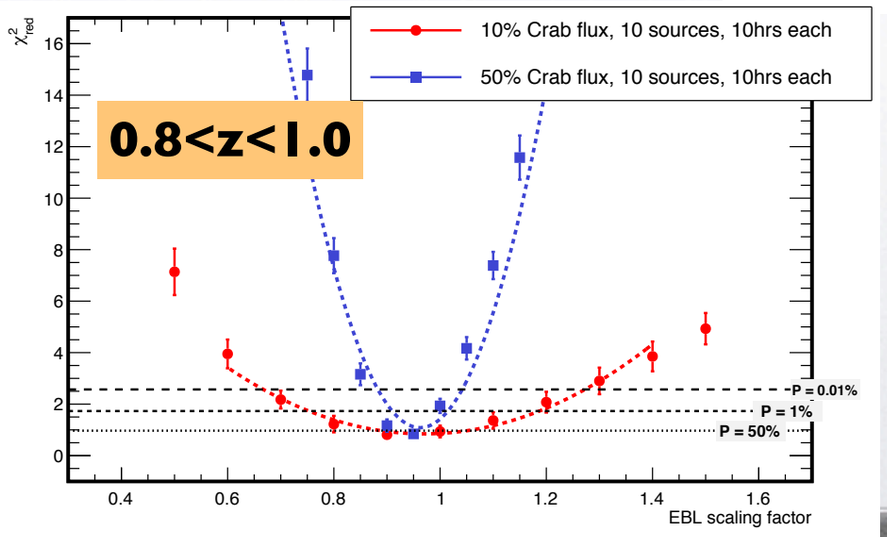
Quality of the fit for different EBL scaling factors





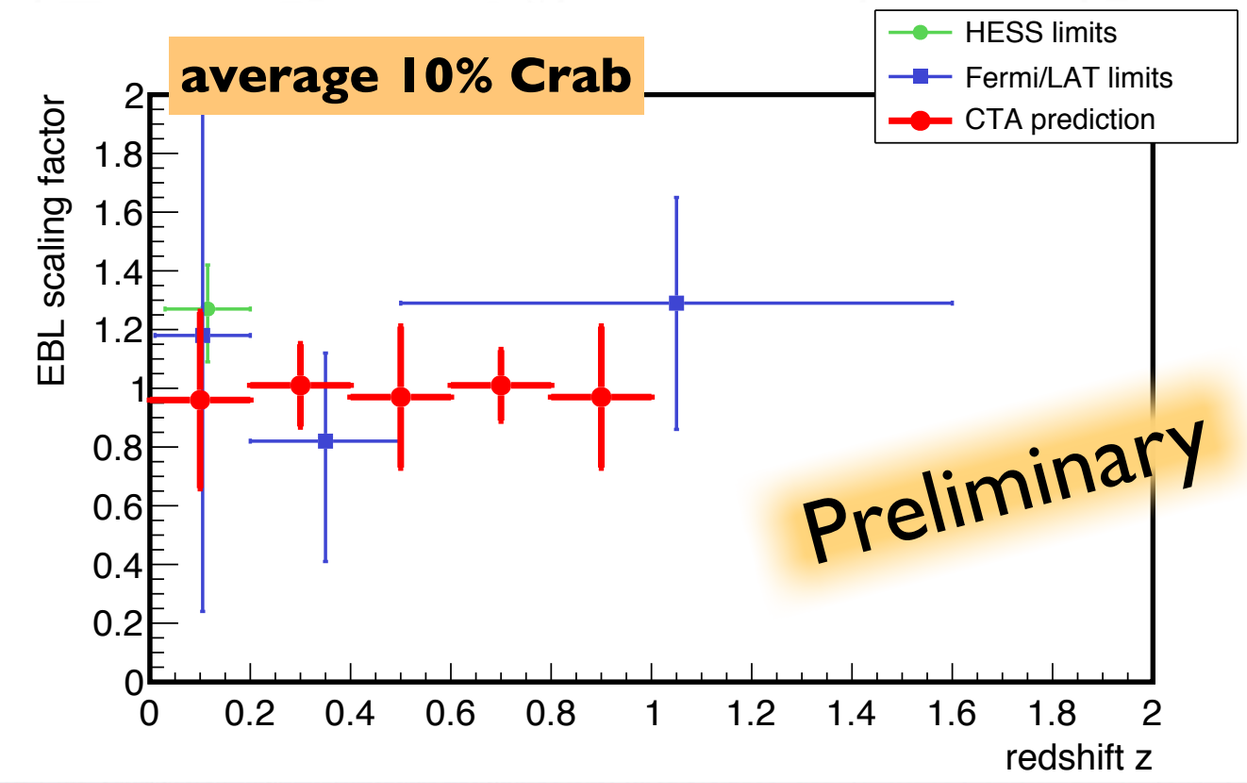


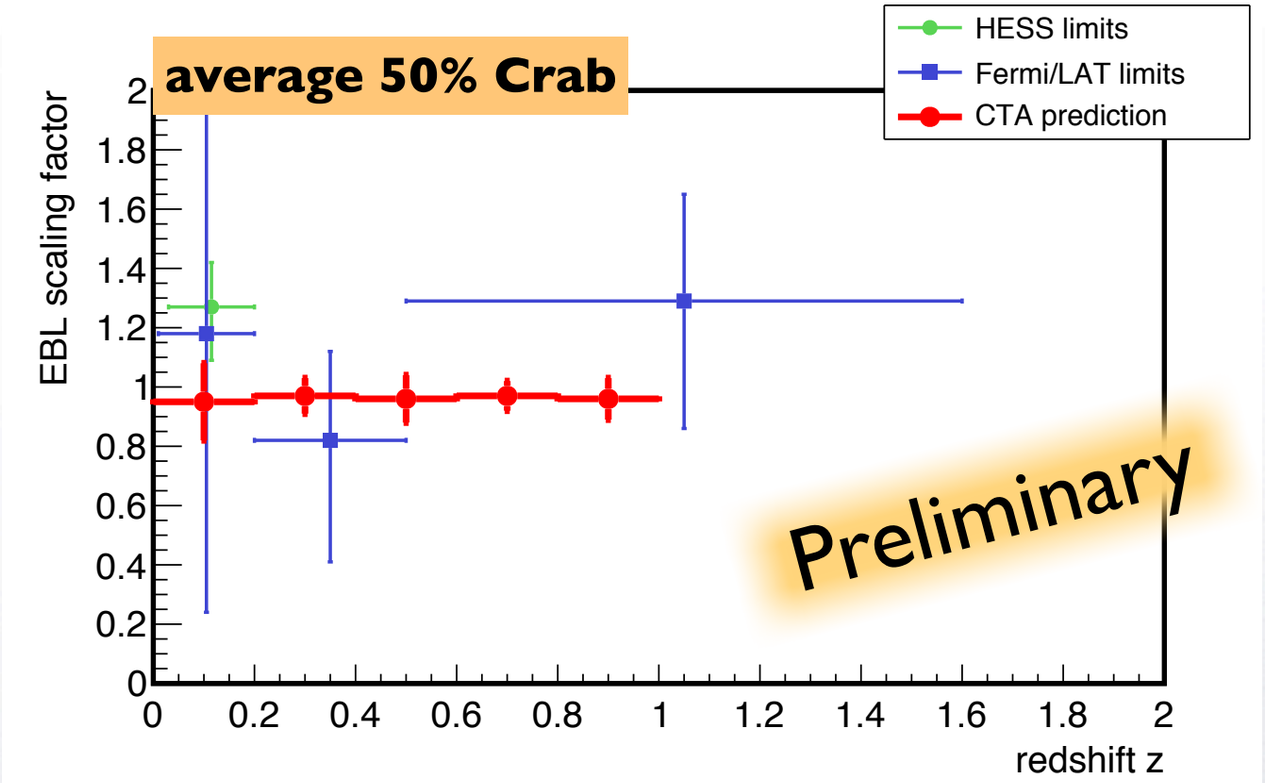




Precision of EBL determination (update)

- Obtained precision (preliminary) of the EBL level: 15-25% for 10% Crab level flares and and 10-15% for the 50% Crab flares
- Can go beyond z>1 if there are sources detectable by CTA





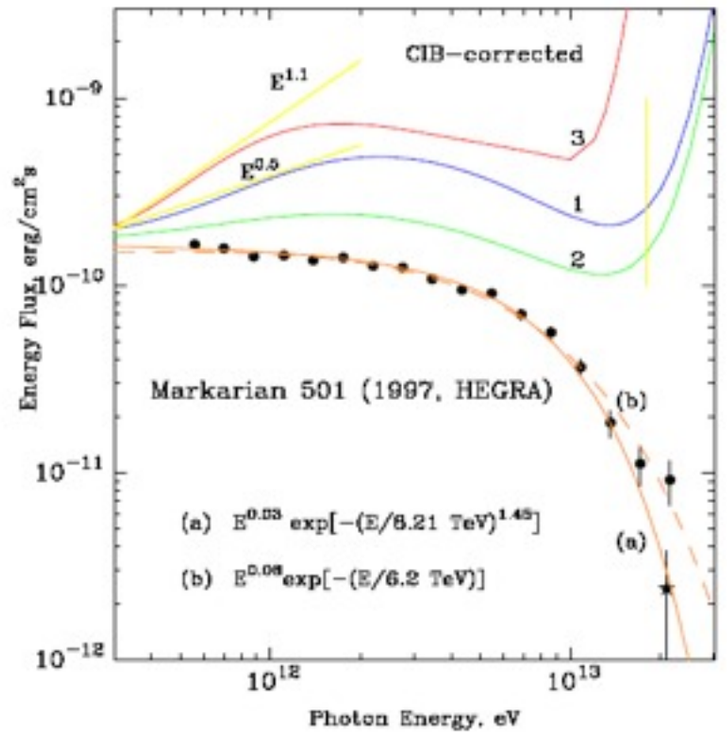
one can derive strong constraints as long as one has enough sources

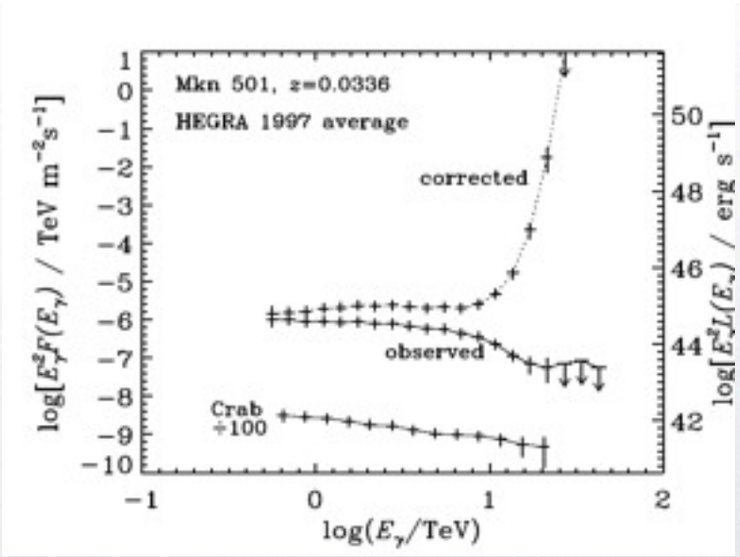
Pile ups in spectra of distant sources

can we actually assume smooth spectra to derive EBL constraints?

Earlier pile-ups

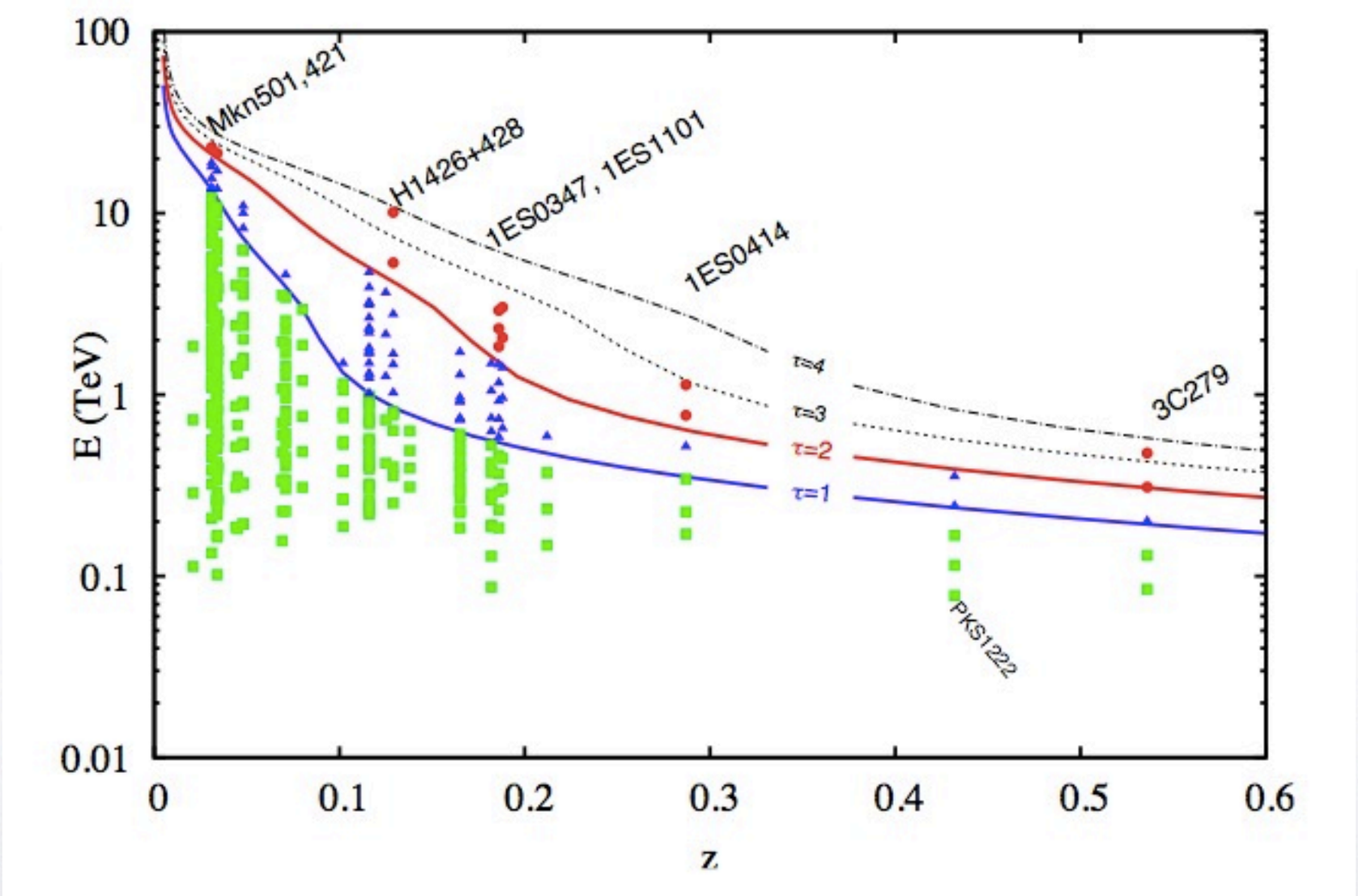
- Earlier observed pile-ups led to:
  - EBL constraints
  - reanalysis of the data
- It is understood that the last spectral point has a strong bias: point derived on a positive fluctuation (no point derived from a negative fluctuation)





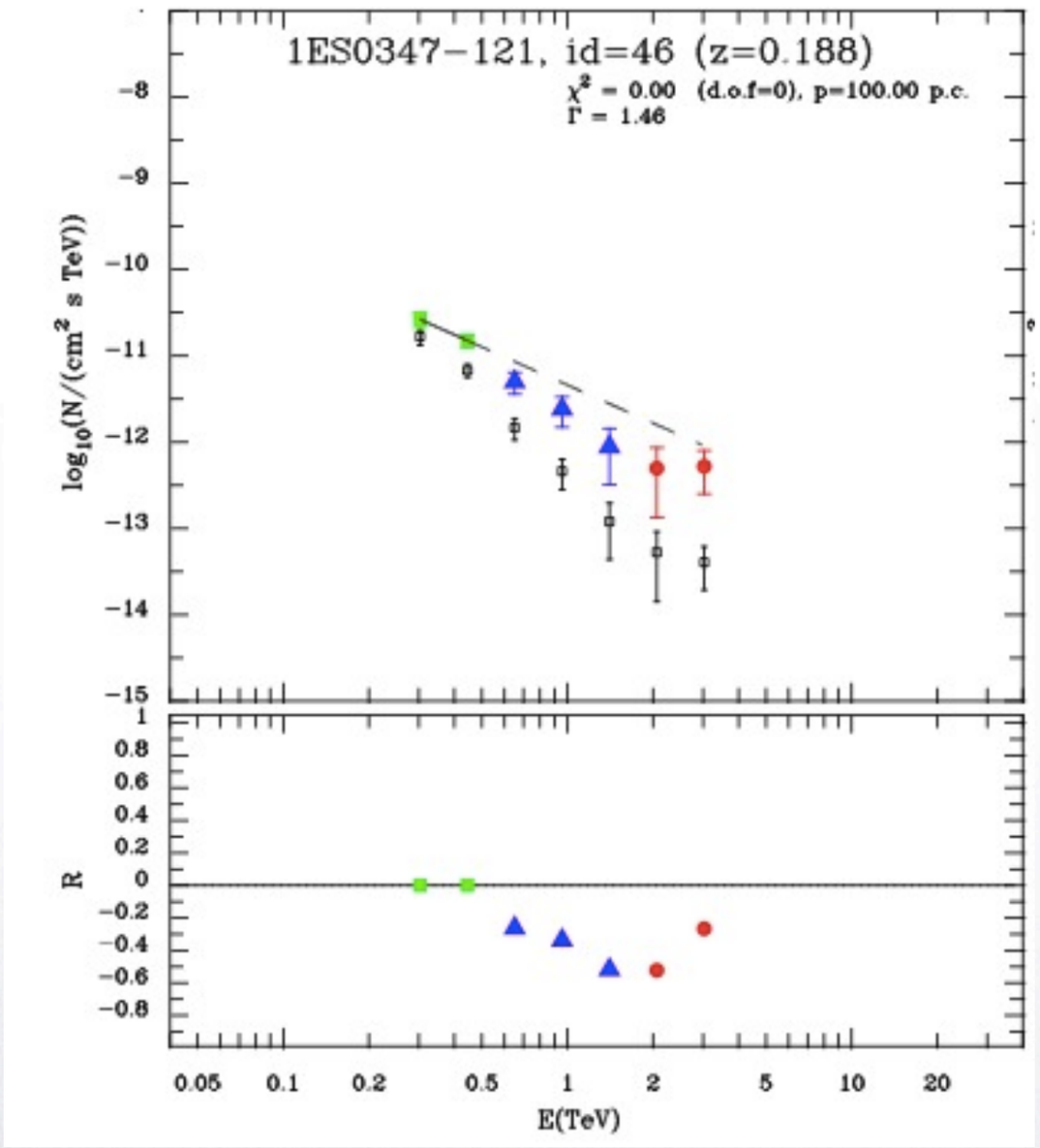
More recent hints

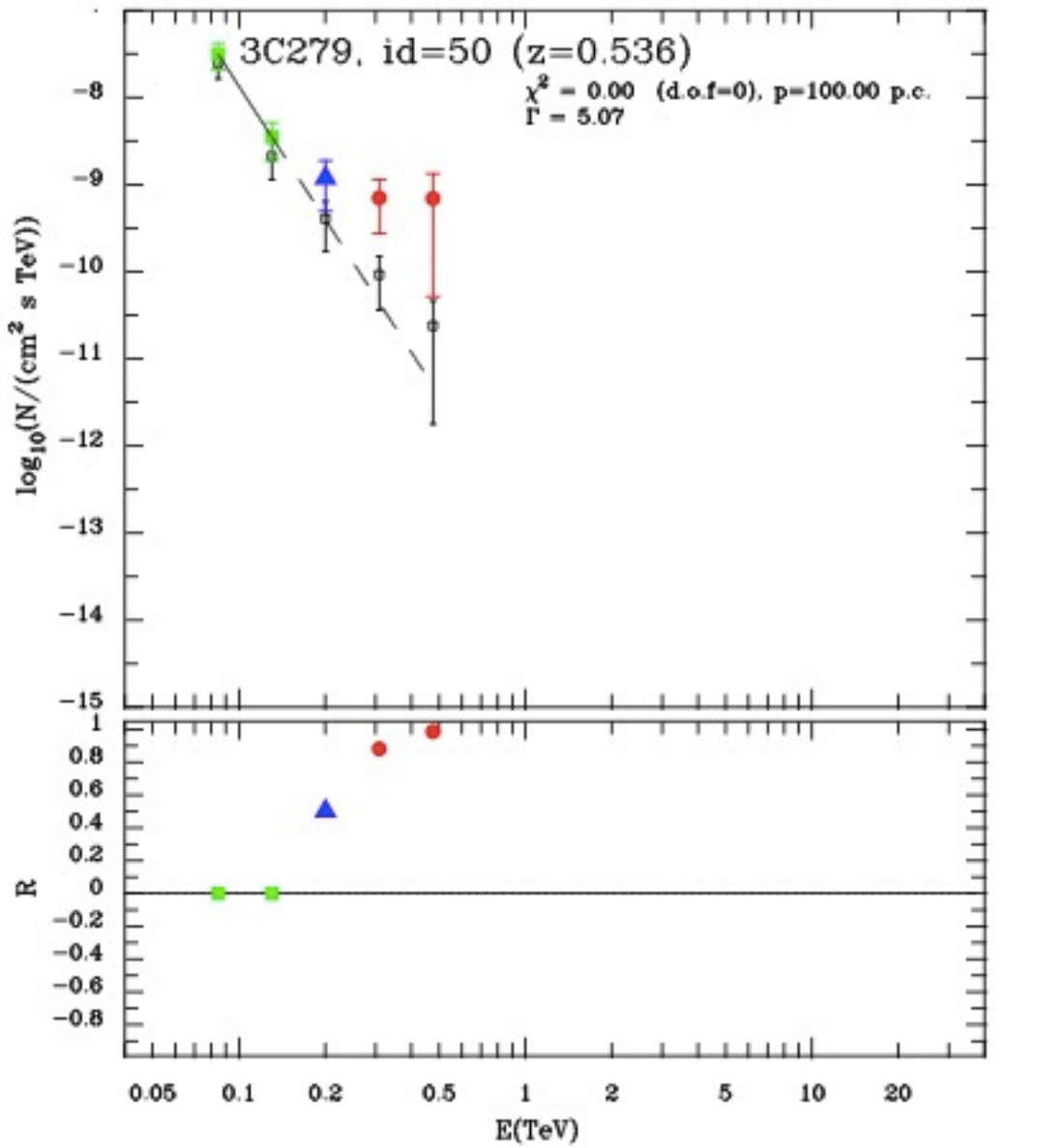
highest points in measured VHE spectra often have tau>2



Horns&Meyer 2012

Example spectra

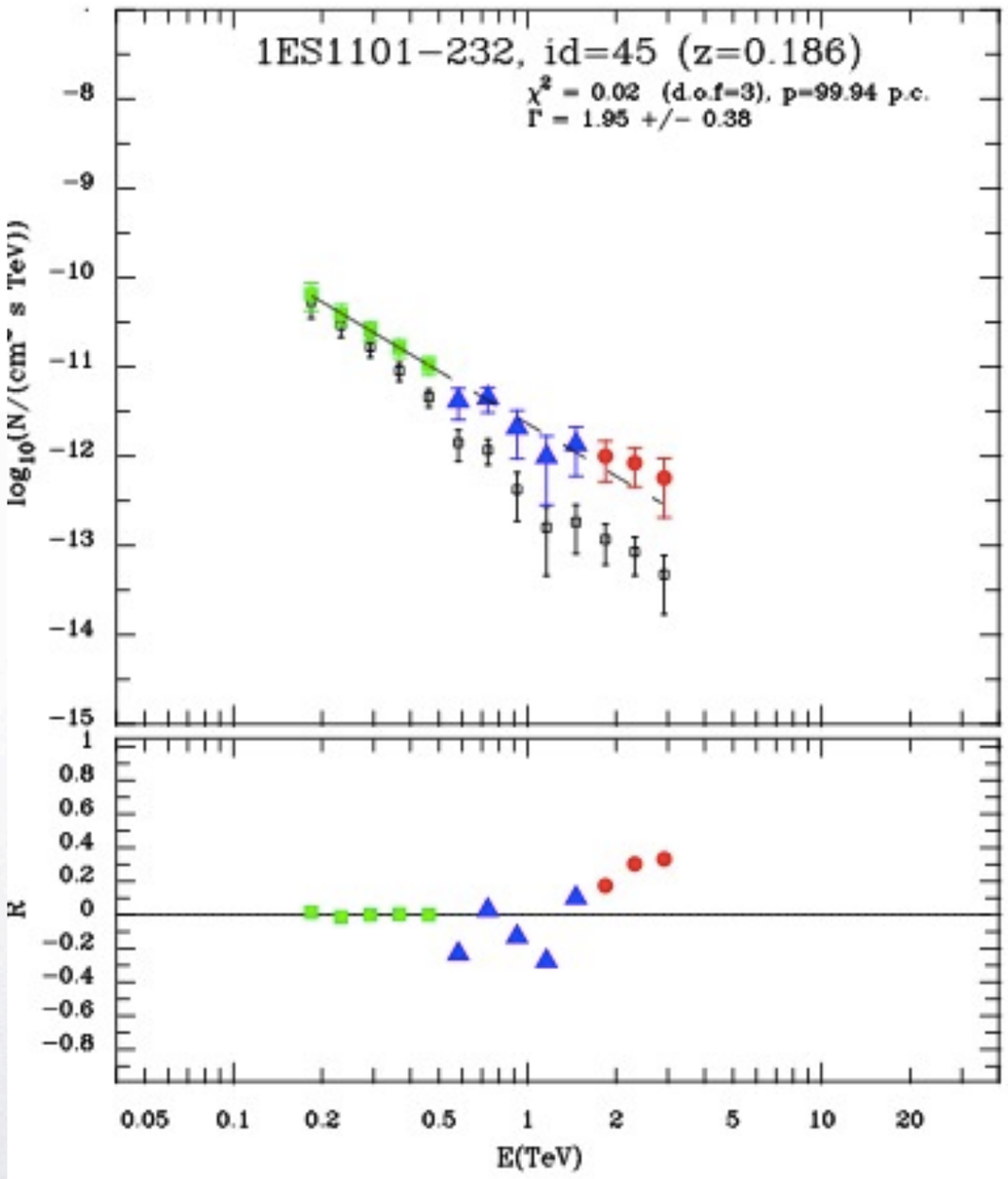


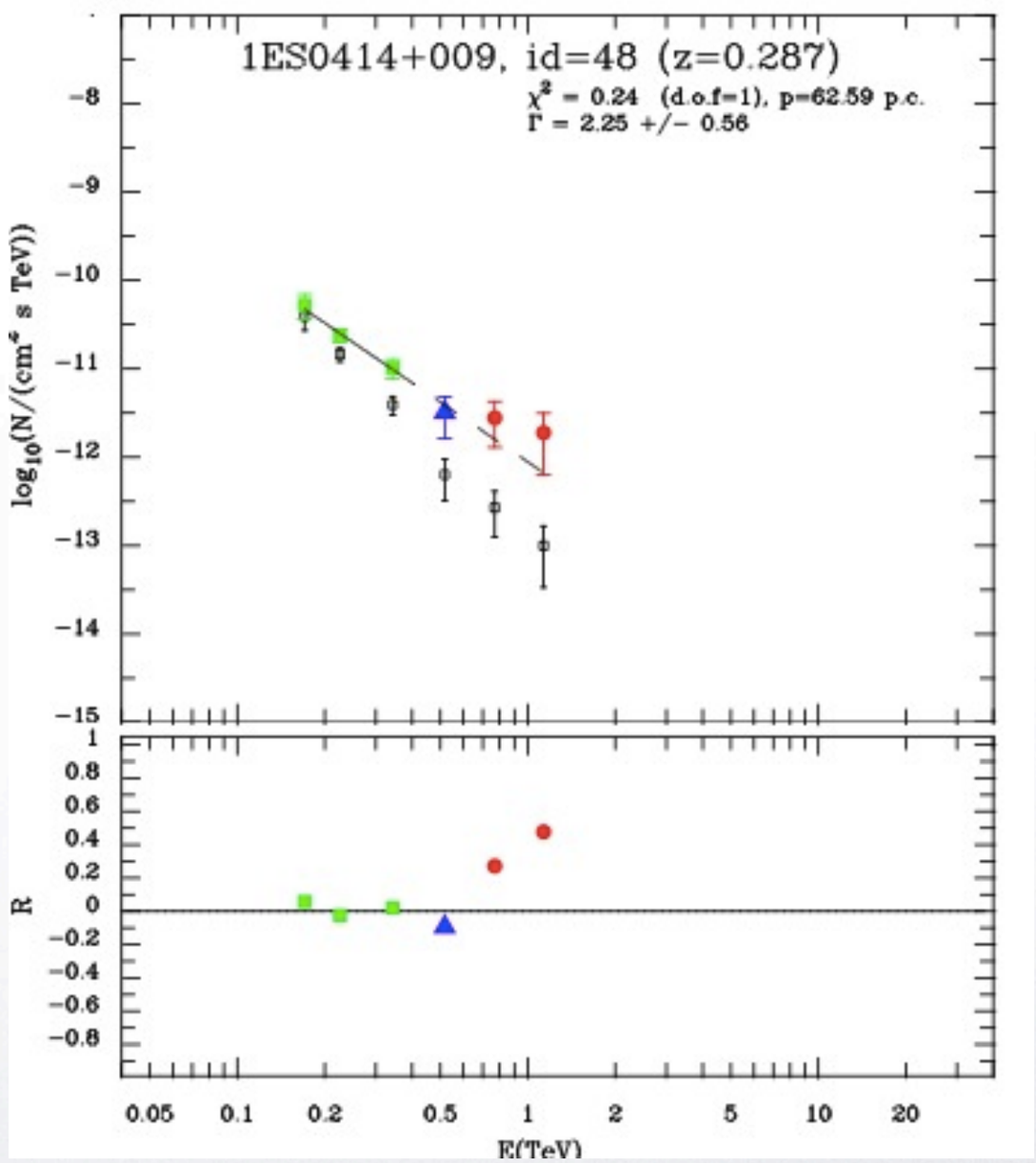


blue: 1 < τ < 2

red: τ > 2

Example spectra



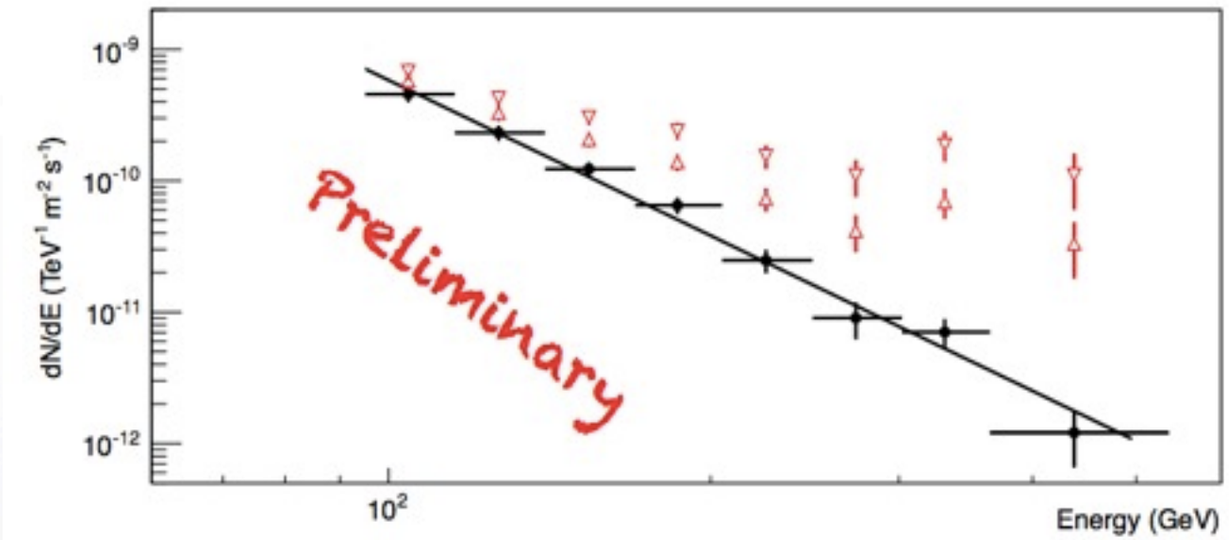


blue: 1 < τ < 2

red: τ > 2

Latest pile-ups: quite significant!

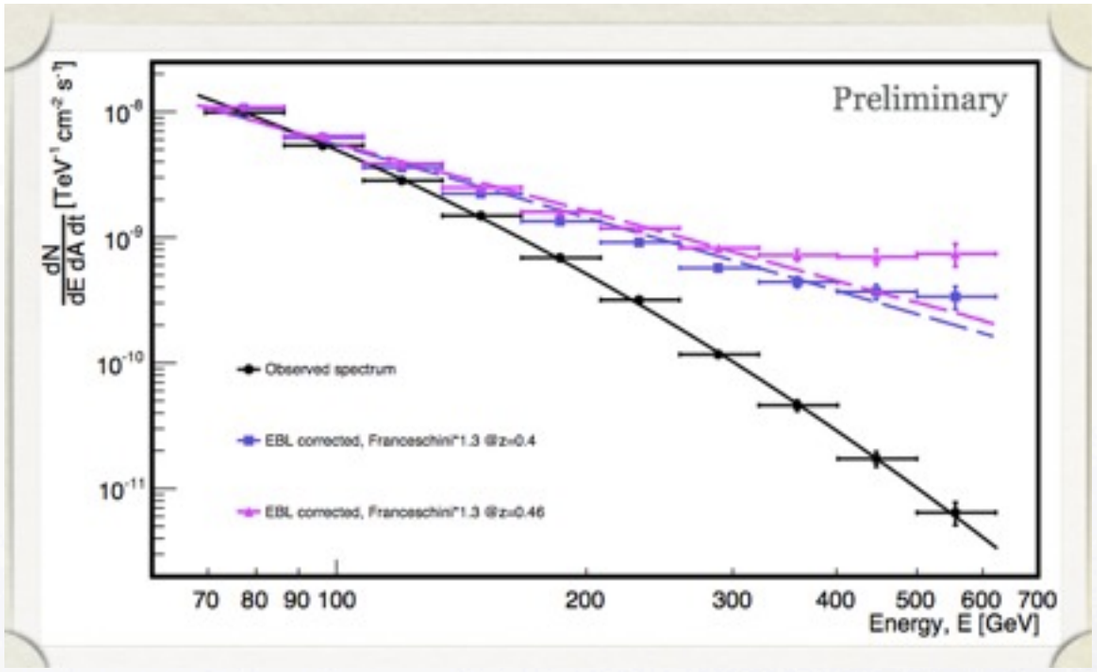
Amy Furniss et al., VERITAS
PKS1424+240, z>0.6



- Gilmore: Γ=2.4±0.2
- Finke: Γ=1.7±0.2

Neither is strictly constraining by Γ>1.5, but the spectral shape starts to curve upward above 300 GeV with even the lowest density EBL models

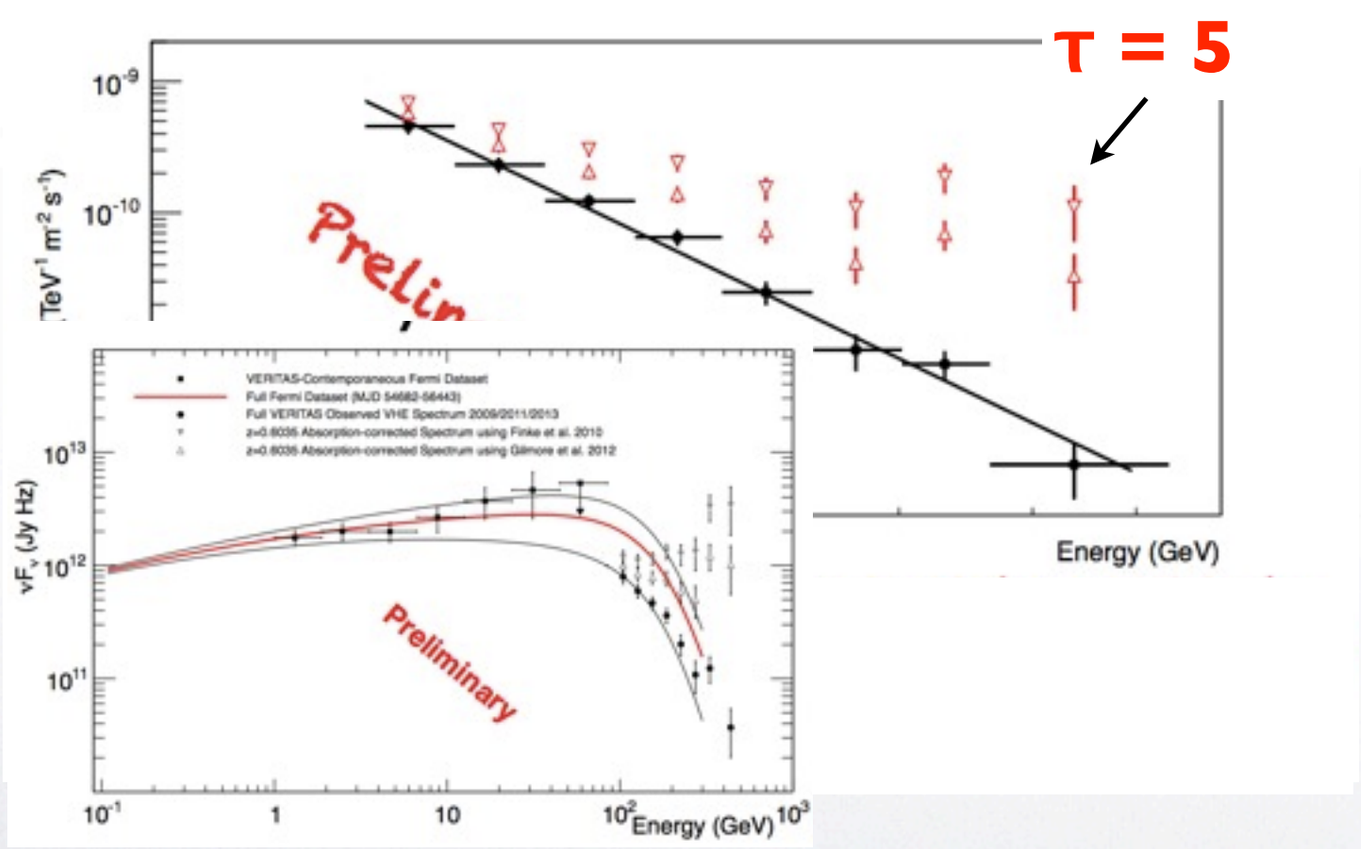
Pepa Becerra et al., MAGIC
PG1553+113, z>0.4



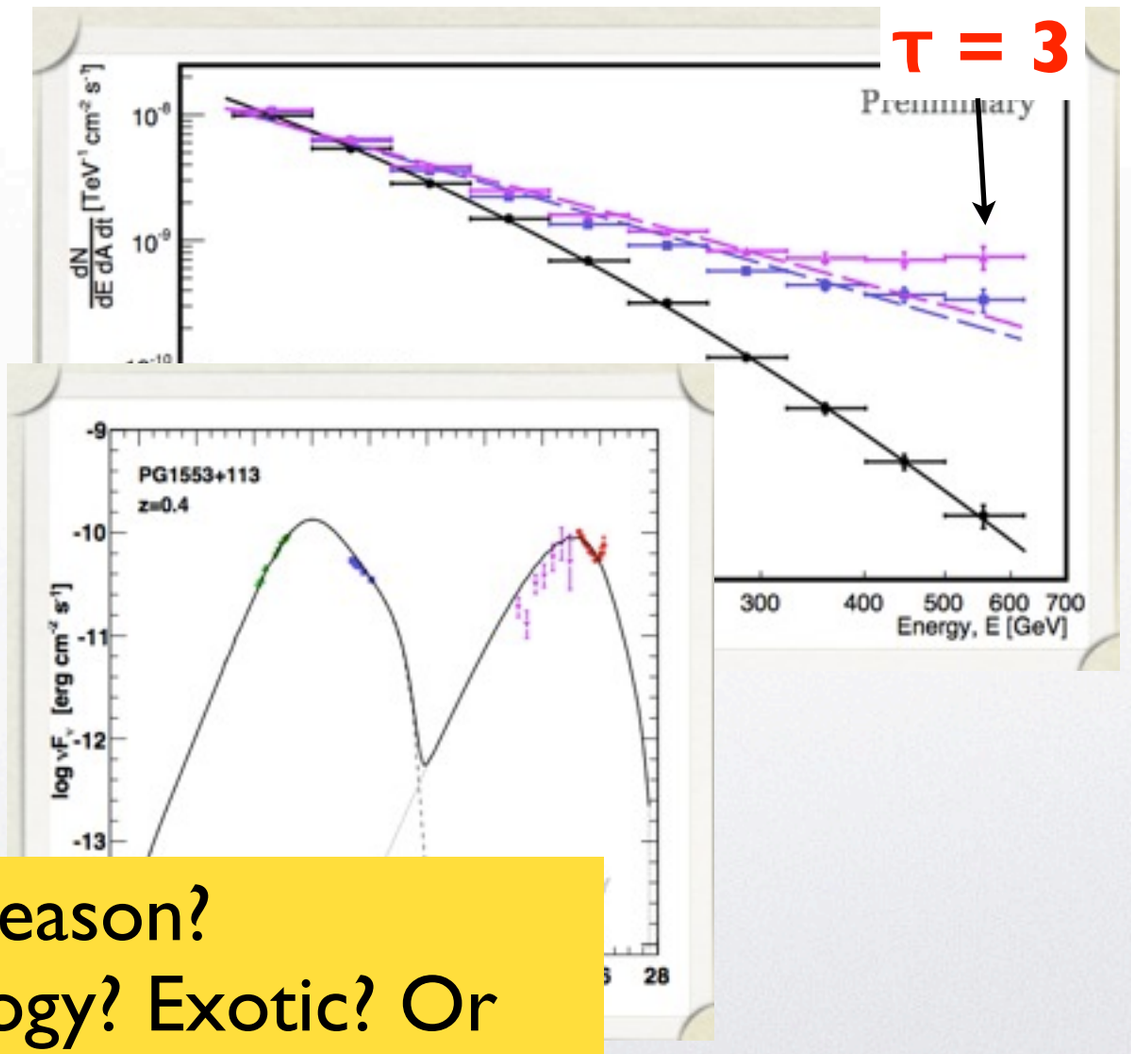
we cannot go lower in the EBL density!

Pile-ups at high energies

Amy Furniss et al., VERITAS
PKS1424+240, z>0.6



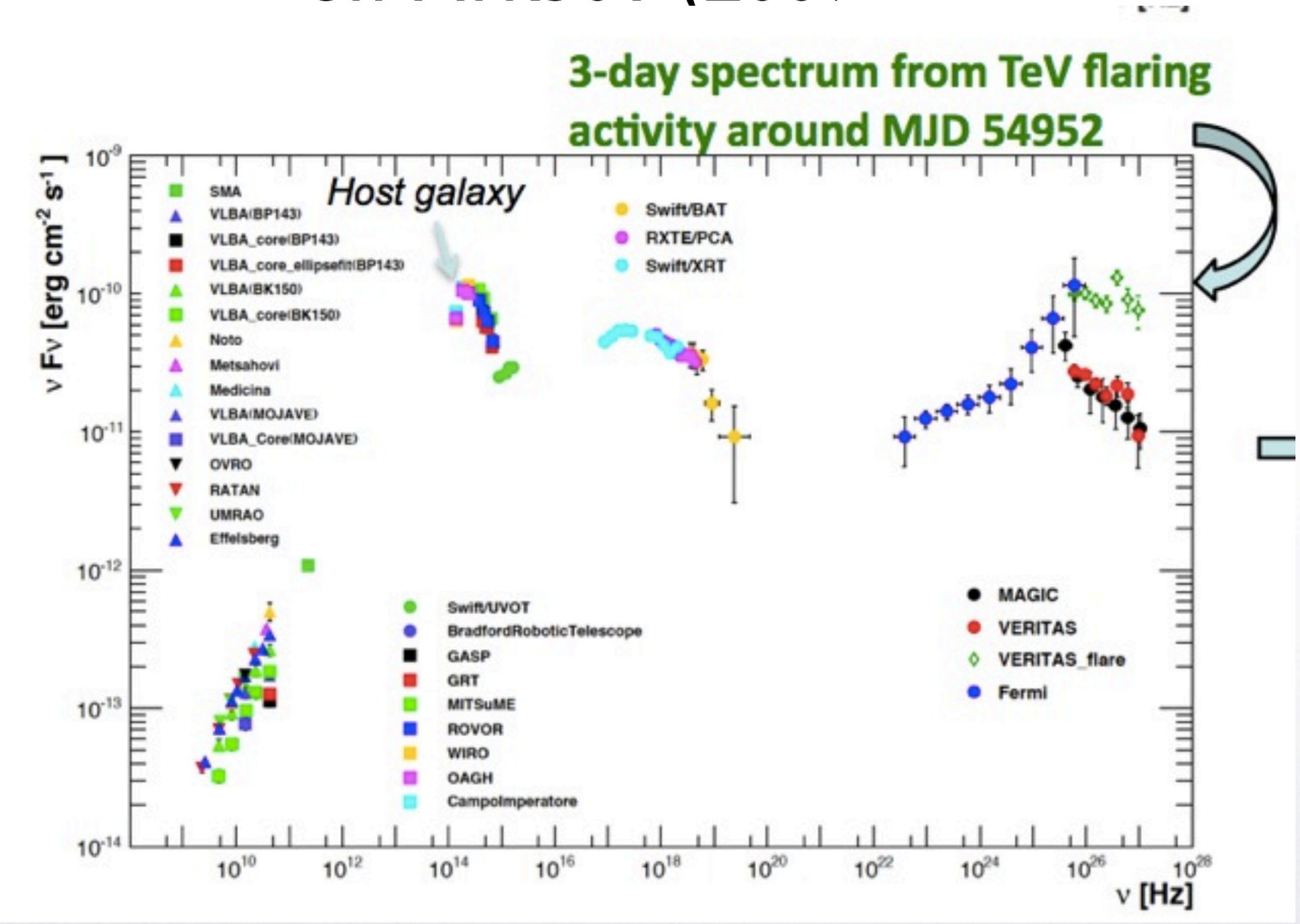
Pepa Becerra et al., MAGIC
PG1553+113, z>0.4



What is the reason?
Energy scale? Cosmology? Exotic? Or simply averaging different flux states?

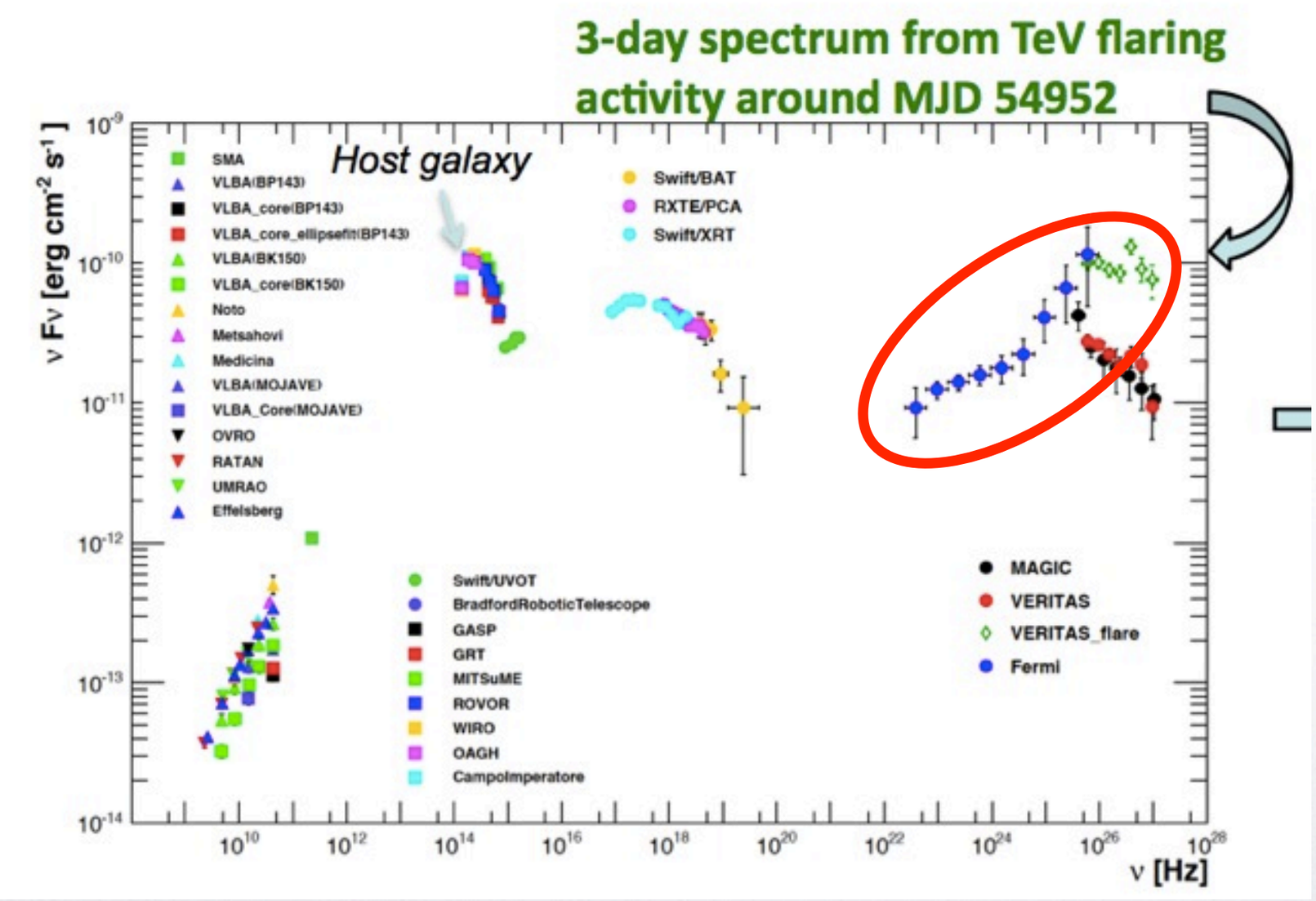
Averaging spectra from different flux states

- example from Fermi/LAT on Mrk501 (2009

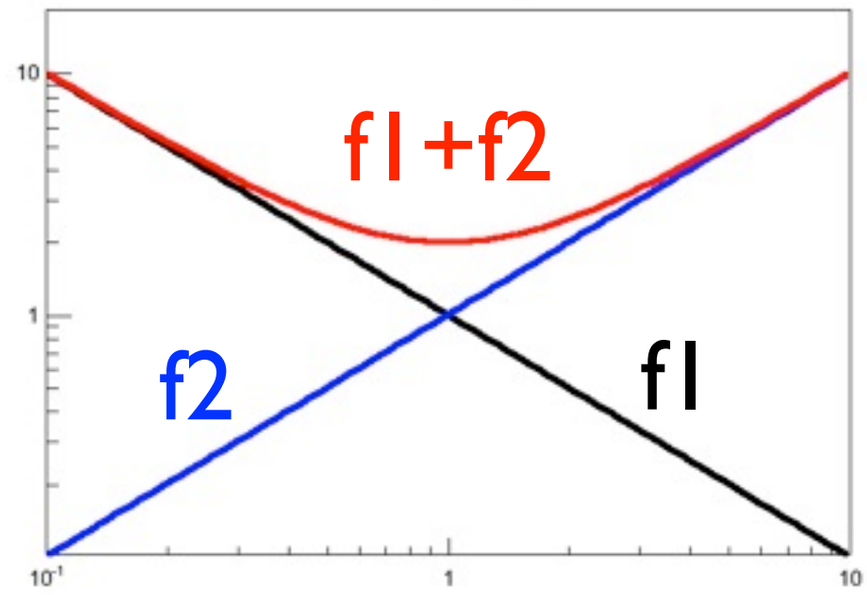


Averaging spectra from different flux states

- example from Fermi/LAT on Mrk501 (2009



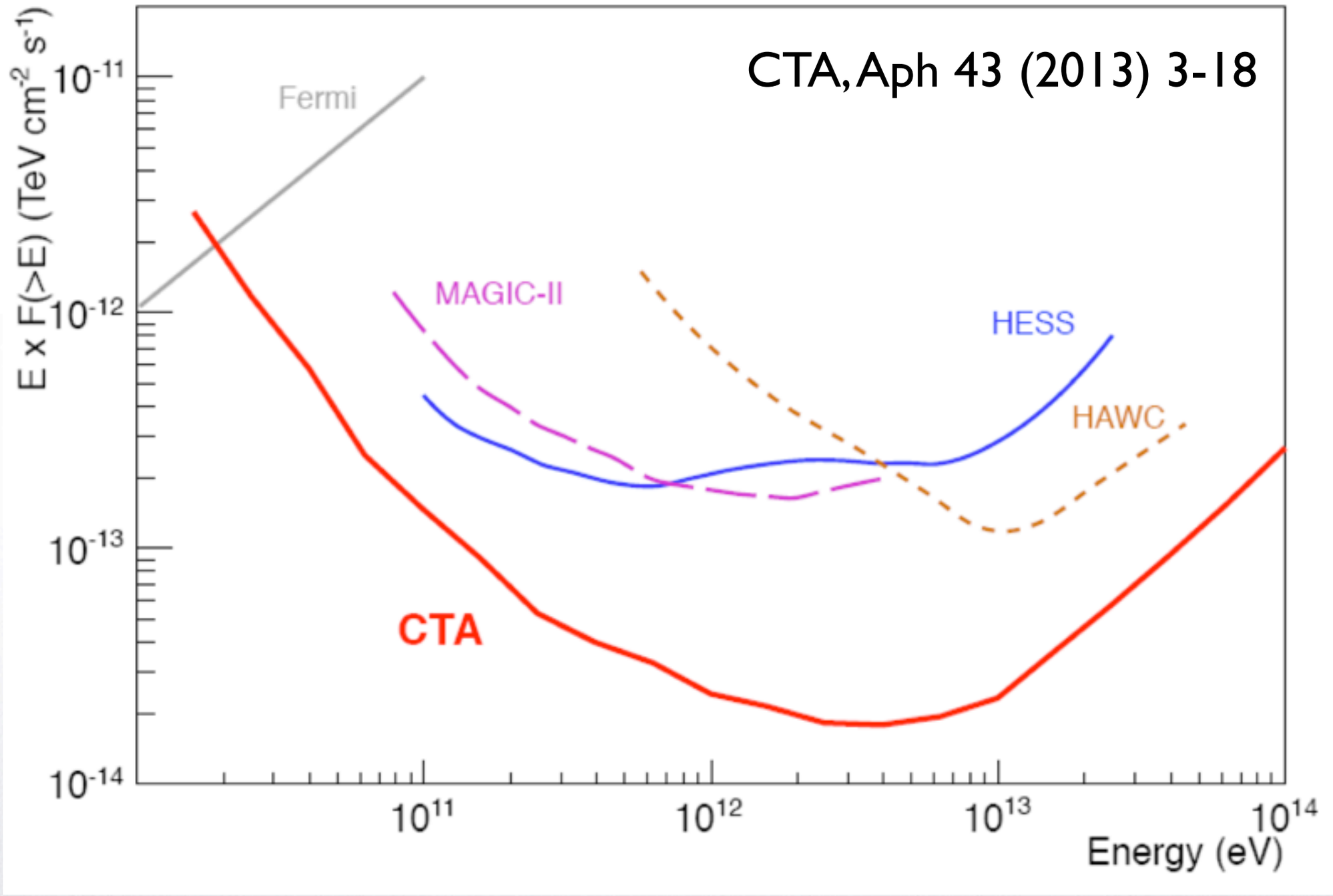
- mixing flux states



need to be careful when interpreting averaged spectra

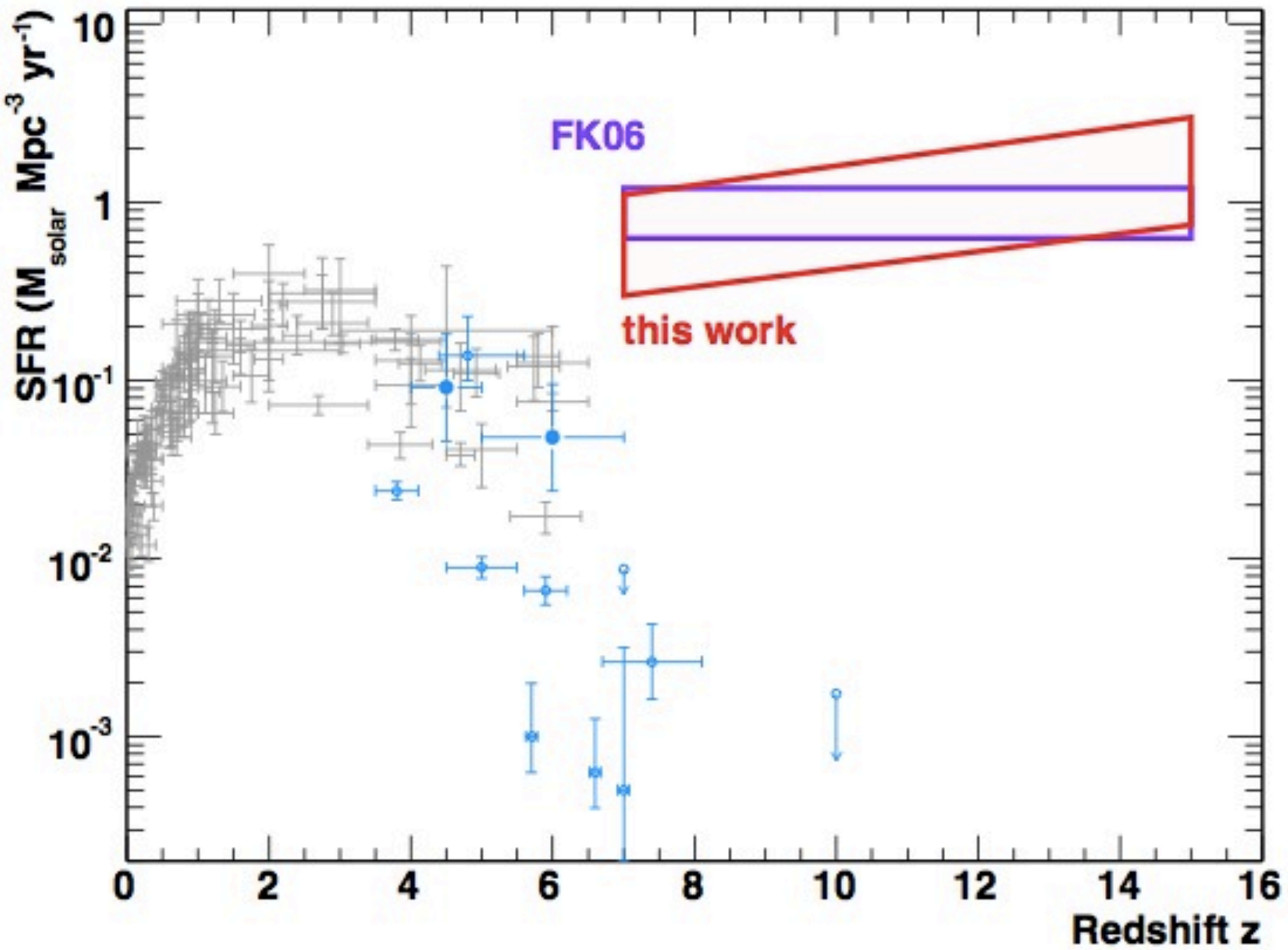
Sources of the EBL

Sensitivity



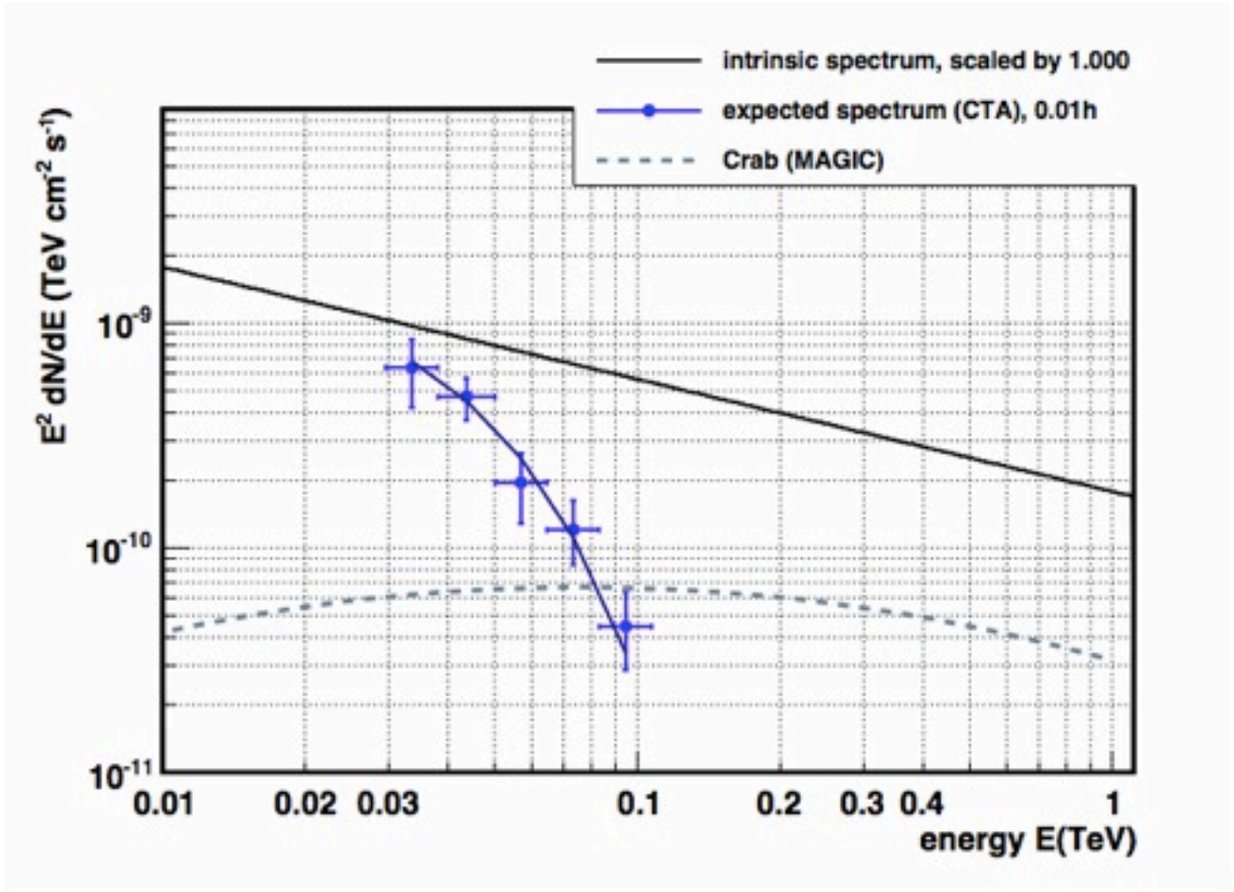
Beacons beyond the peak in the star formation rate

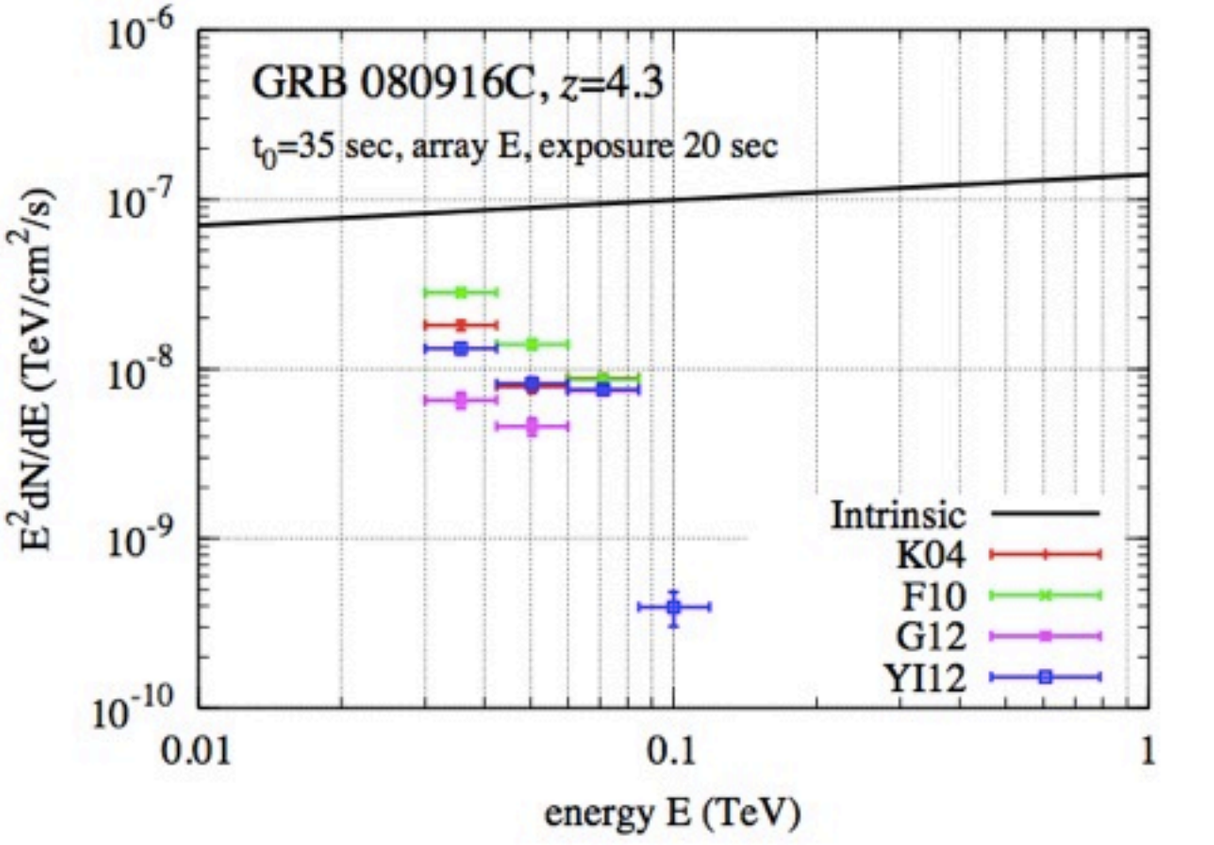
Raue, Kneiske, DM, A&A 498, 25 (2009)



GRB case

DM et al. AP, 43 (2013), 241



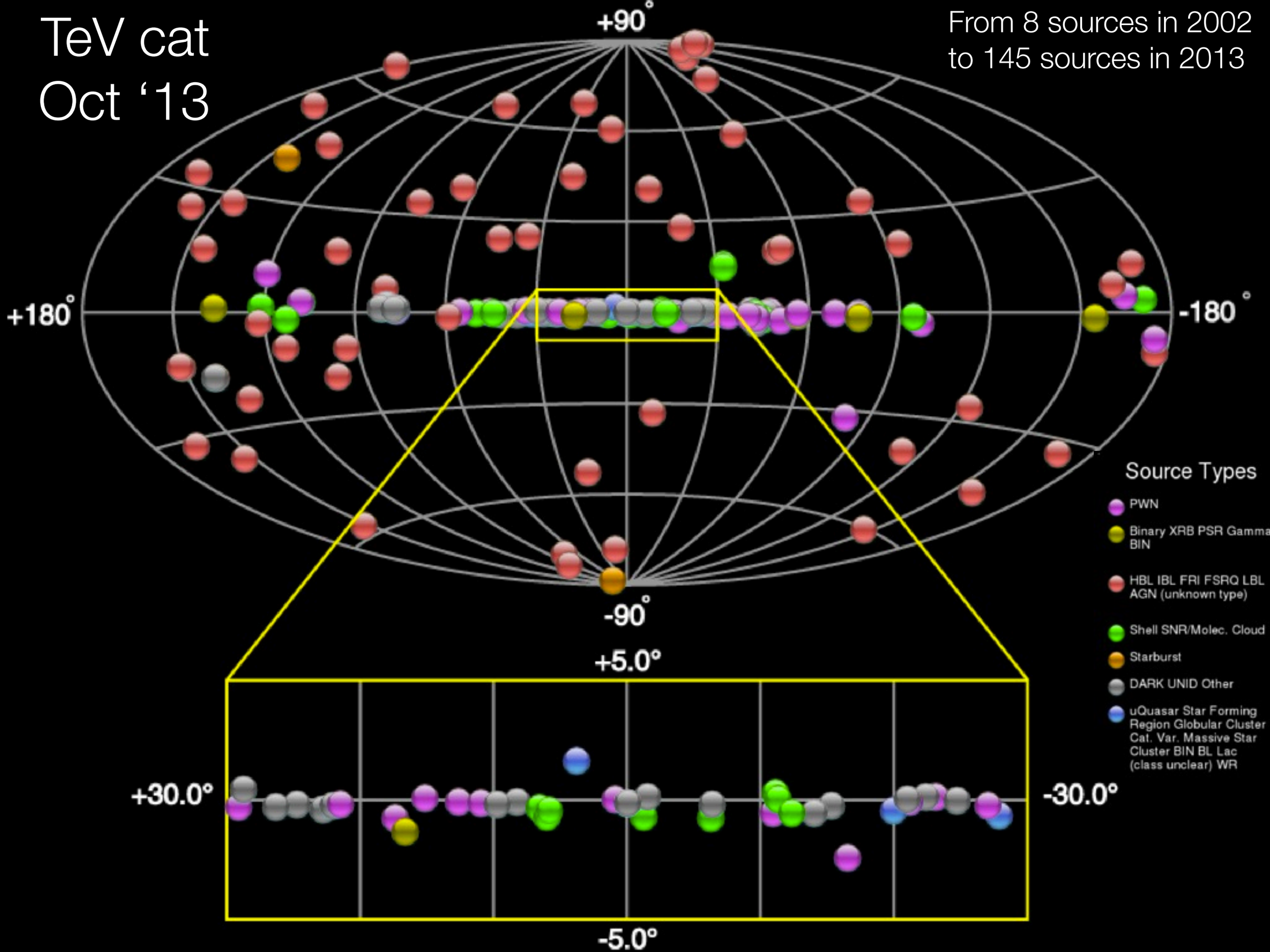


Simulated energy spectrum of GRB 080916C (z = 4.3) if measured with CTA.
Left panel: The intrinsic spectrum is assumed to follow $2.9 \times 10^{-9} \times (E/\mathrm{TeV})^{-2.16}[\mathrm{TeV}^{-1}\mathrm{cm}^{-2}\mathrm{s}^{-1}]$ and the duration of the measurement is 45 s (T0=55–100s, interval "e") as measured with Fermi.
Right panel: Exposure time 20 sec for the interval "d", i.e. the assumed source flux is $dN/dE = 1.4 \times 10^{-7}(E/\mathrm{TeV})^{-1.85}\mathrm{cm}^{-2}\mathrm{s}^{-1}\mathrm{TeV}^{-1}$.

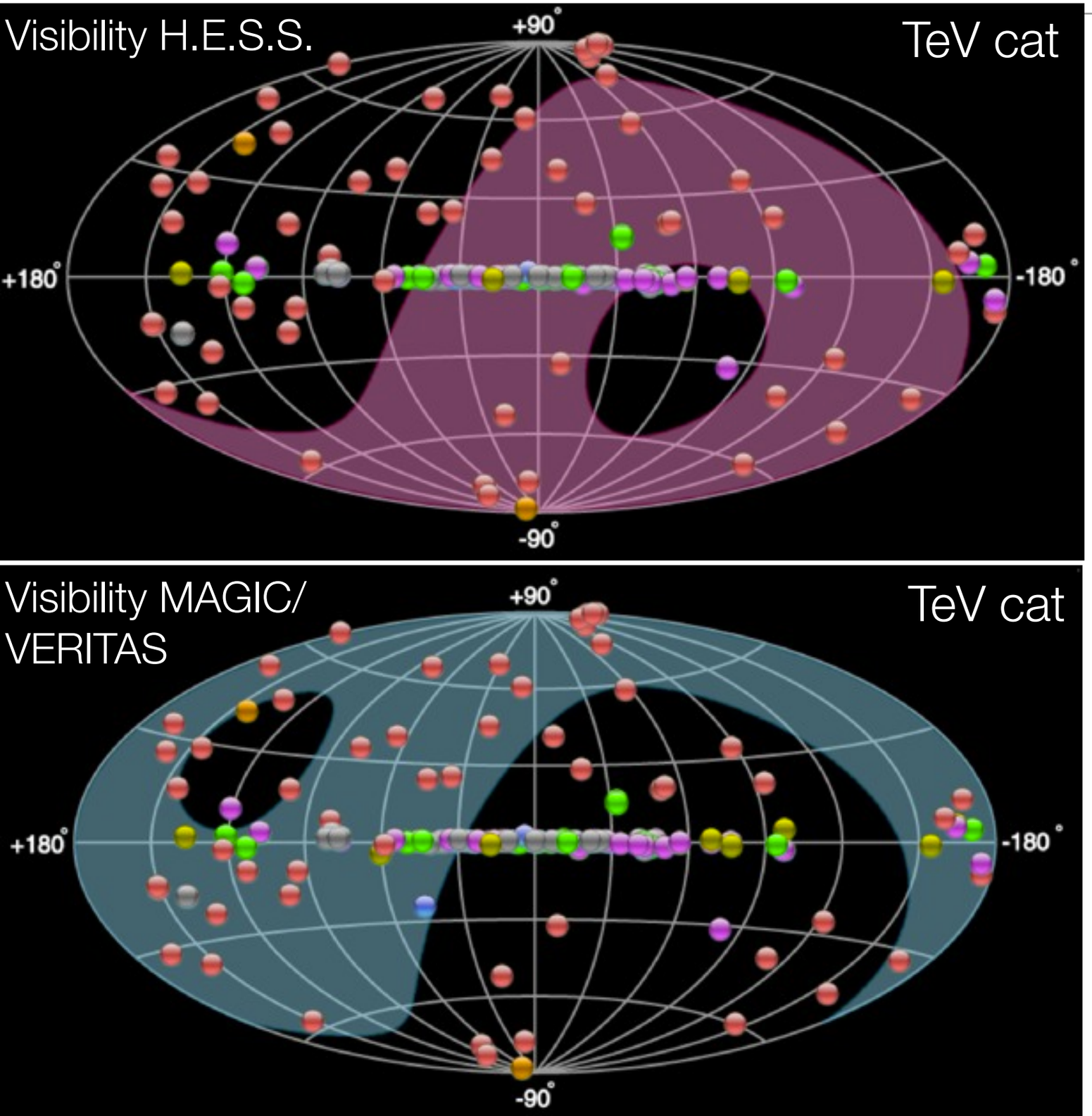
CTA can detect GRBs up to z=6 → HESS-II can do the same!

TeV cat Oct ‘13

From 8 sources in 2002
to 145 sources in 2013



Observability of the sky



- Visibility is shown for culmination below zd=30deg
- One can see that sources with culmination up to zd=45 are also detected (but fewer)
- Good complementarity of the sites
- Number of useful hours: ~1200h per year, including moon time

How many sources will be useful for EBL/cosmology studies?

DM & D. Semikoz 2013, in preparation

- Fermi/LAT catalog of AGN's at E>50 GeV
- Comparison of CTA sites for detection of Fermi AGN's
- Detection of AGN's within redshift bins, step towards determination of EBL
- Hard spectrum sources and detection of IGMF

Method

DM & D. Semikoz 2013, in preparation

- We used 2nd Fermi LAT catalog of AGN's
- In order to have minimal background in Fermi LAT:
  - We used superclean photons only
  - We separate Galactic plane at |b|=10 degrees
  - We divide sources on BL Lacs and other AGN's

- We used photons with E>50 GeV
- With at least 1 photon in 0.2 degree from sources we have around 10 background photons in the all sky for BL Lacs with |b|>10 degrees
- Flux of sources is averaged over 5 years of Fermi
- We used Fermi power law spectra cutoff at 1 TeV

Method

DM & D. Semikoz 2013, in preparation

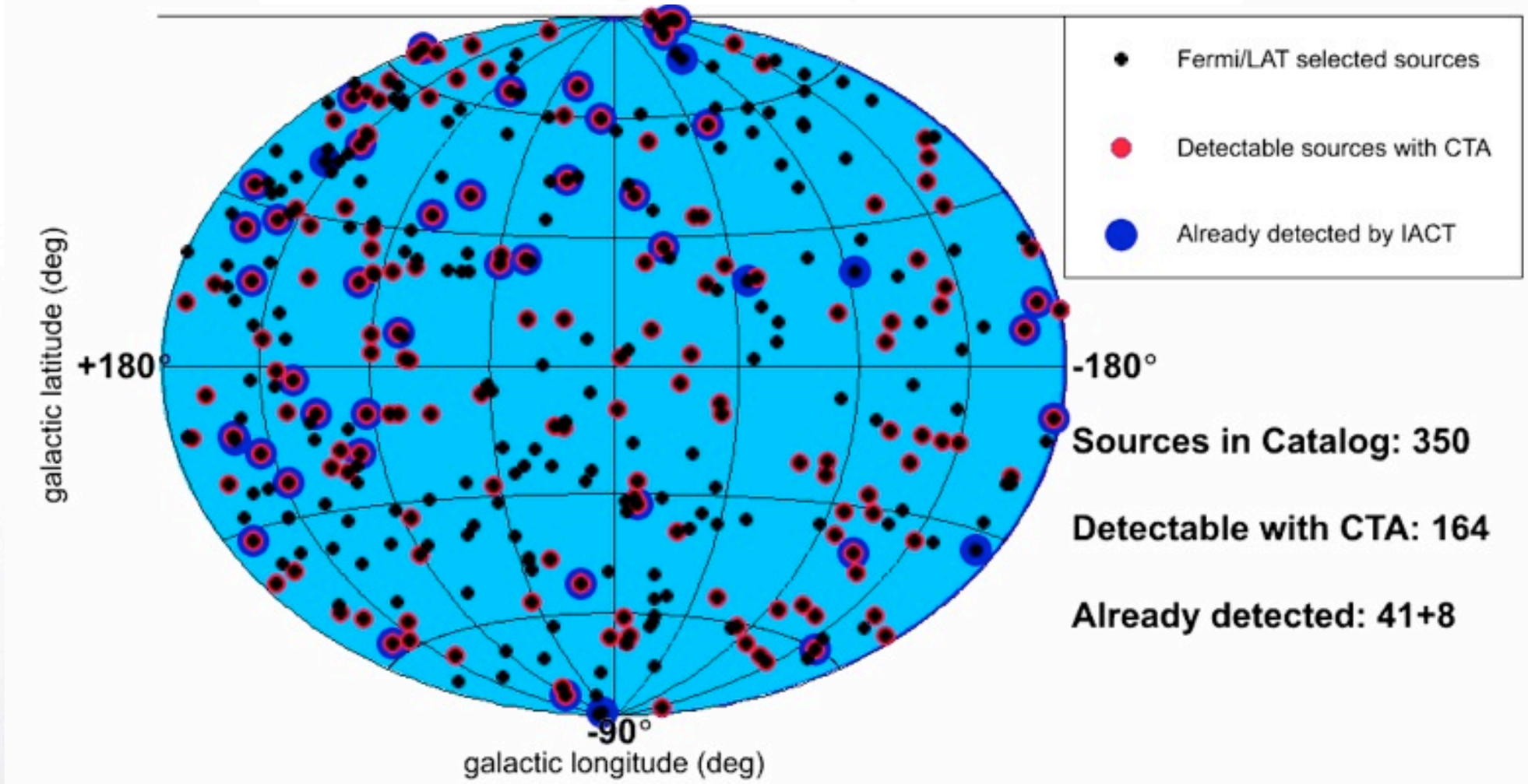
- Fermi/LAT spectra are extrapolated to the CTA energy regime using the redshift when available (sources with no redshift are skipped)
- Exposure of 20h (50h) per source
- EBL model of Franceschini et al. (2008) is used
- Only sources that culminate below 30 (45) deg at a given site are "detectable"
- Caution: these are averaged fluxes, no flaring activity is considered in the analysis

Results

DM & D. Semikoz 2013, in preparation

20h exposure of every source

CTA generic everywhere



1) from 350 hard spectra sources, CTA can detect 164 (47%) if built everywhere

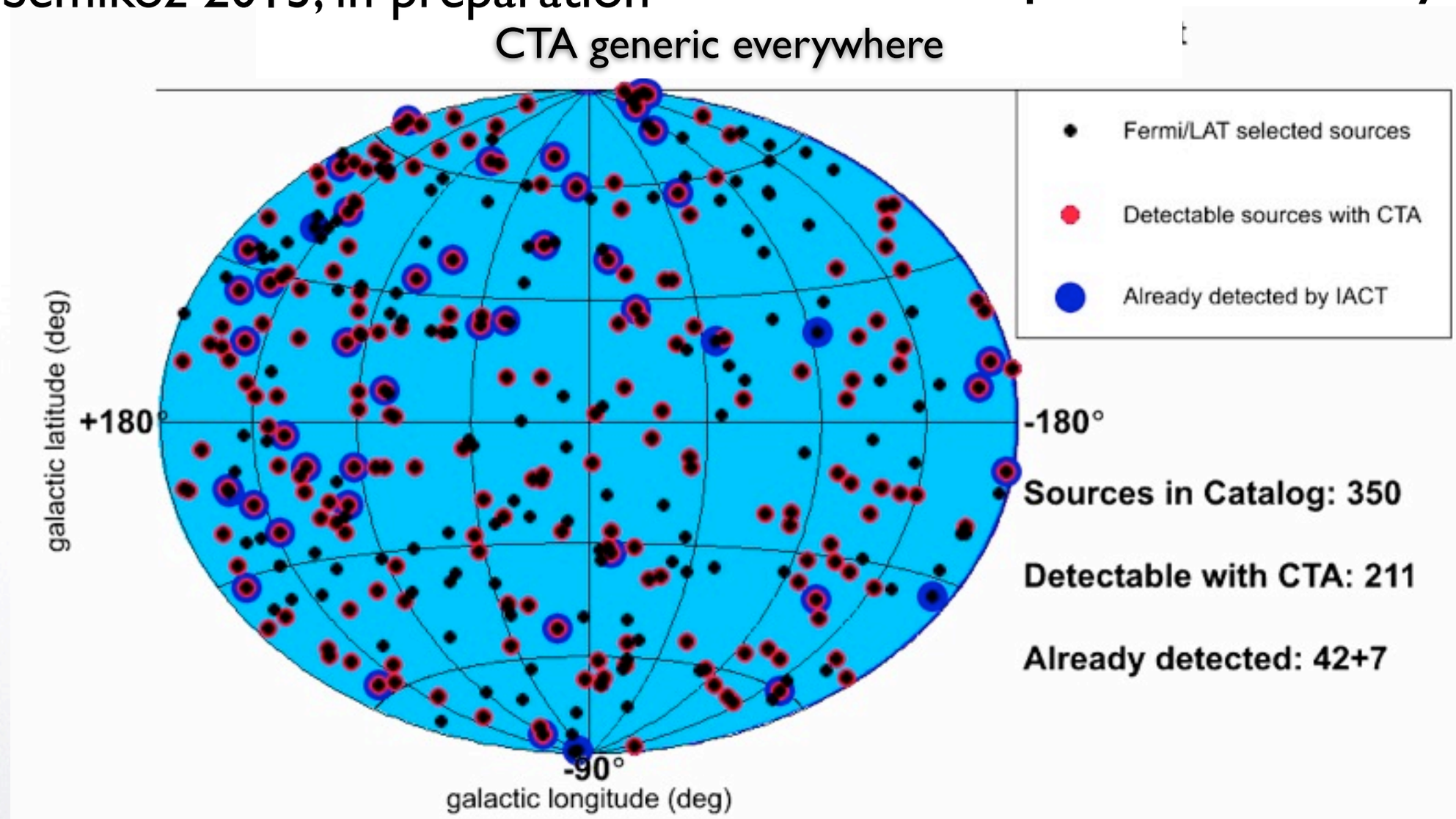
2) 49 sources (14%) already detected

Results

DM & D. Semikoz 2013, in preparation

50h exposure of every source

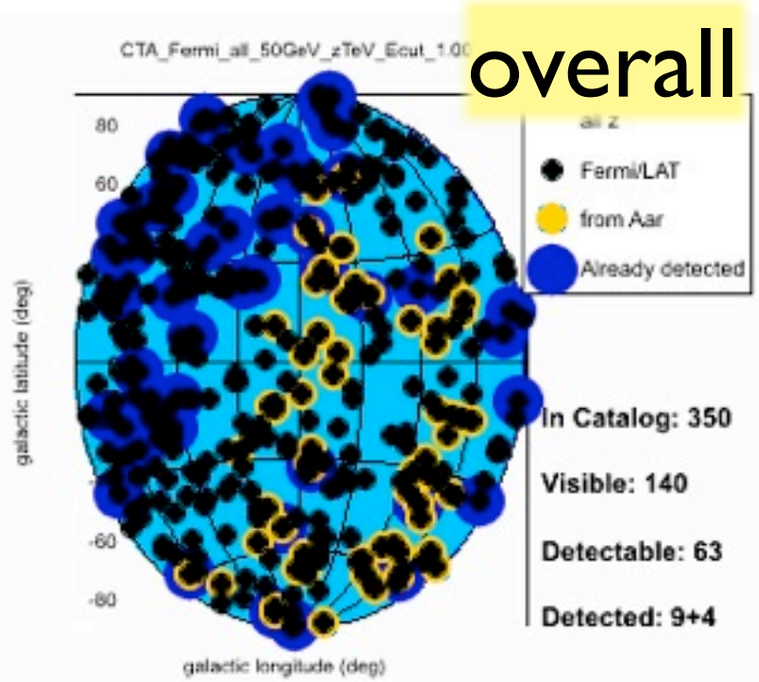
CTA generic everywhere

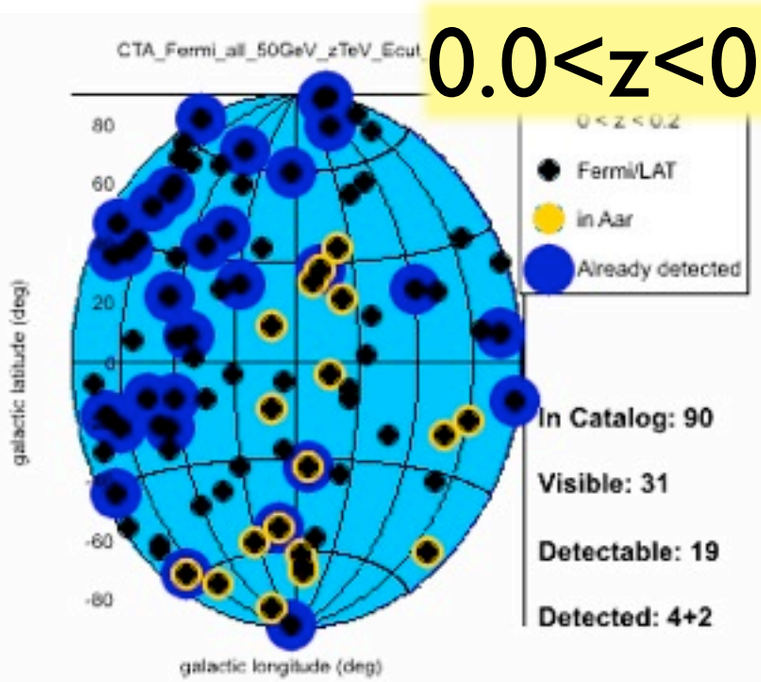


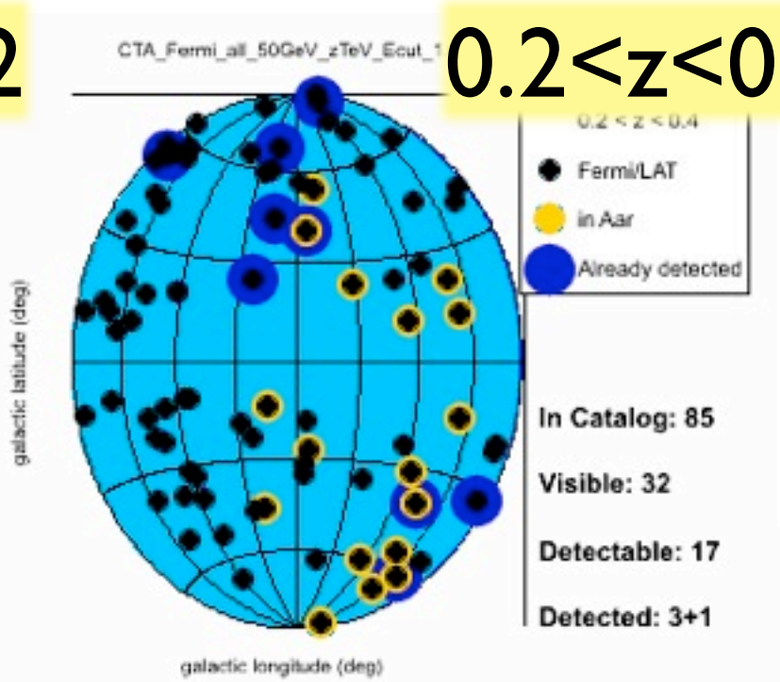
1) from 350 hard spectra sources, CTA can detect 211 (60%) if built everywhere

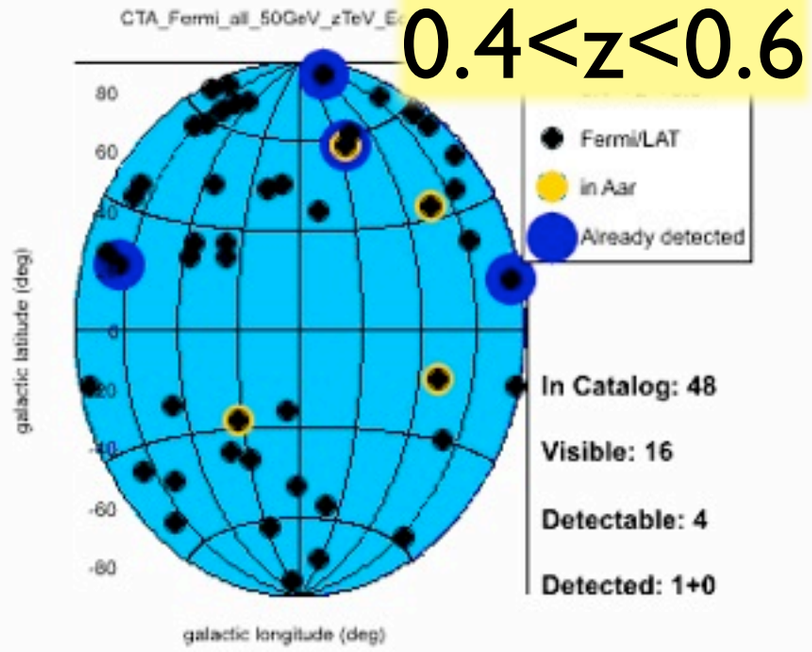
2) 49 sources (14%) already detected

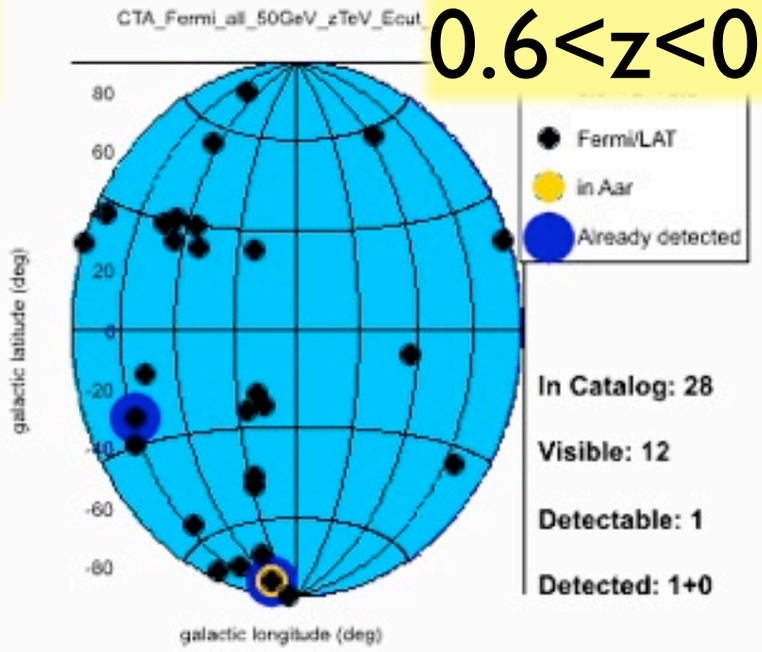
Expectation: CTA in South (20h)

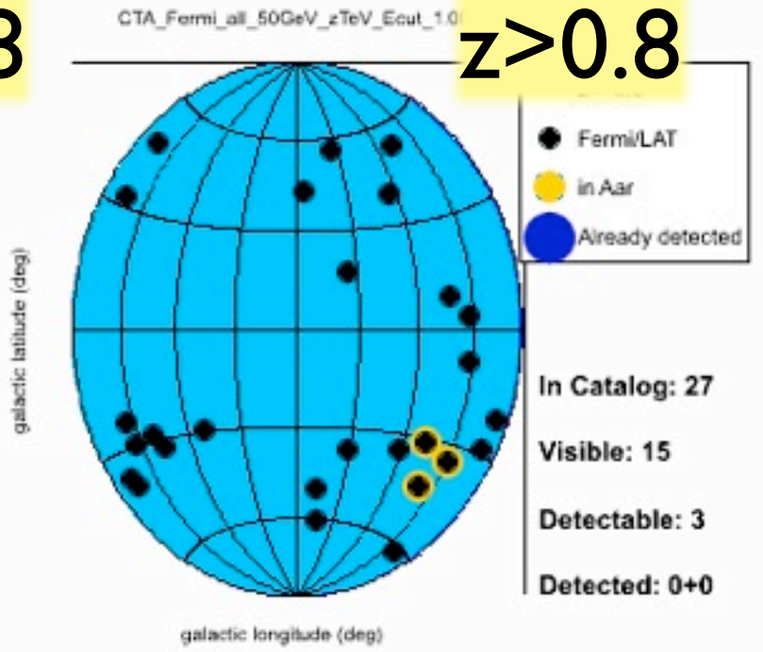






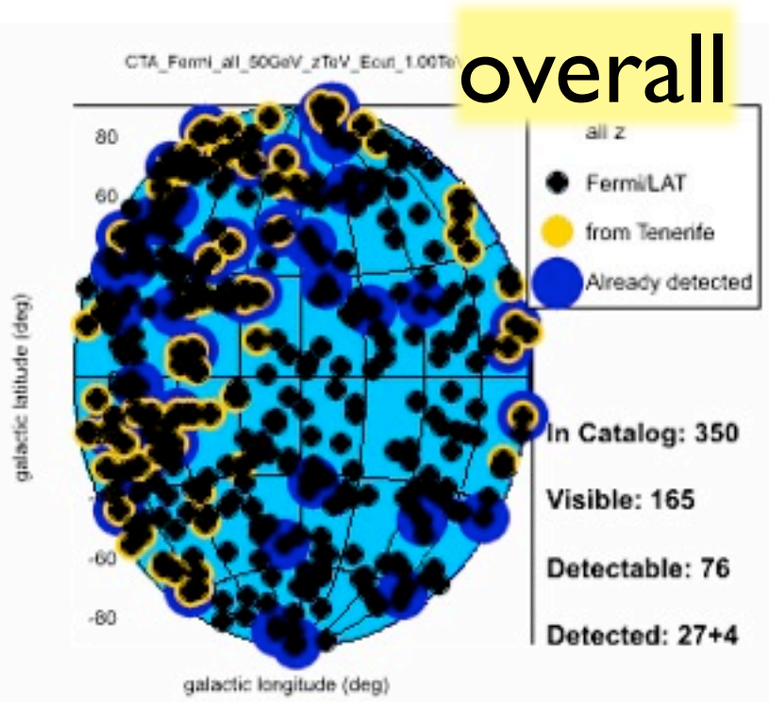


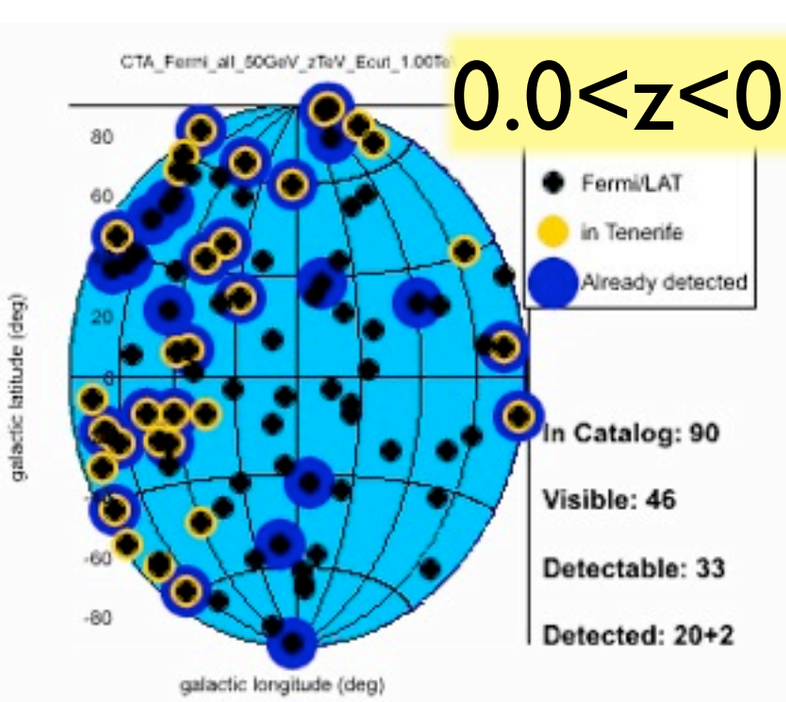


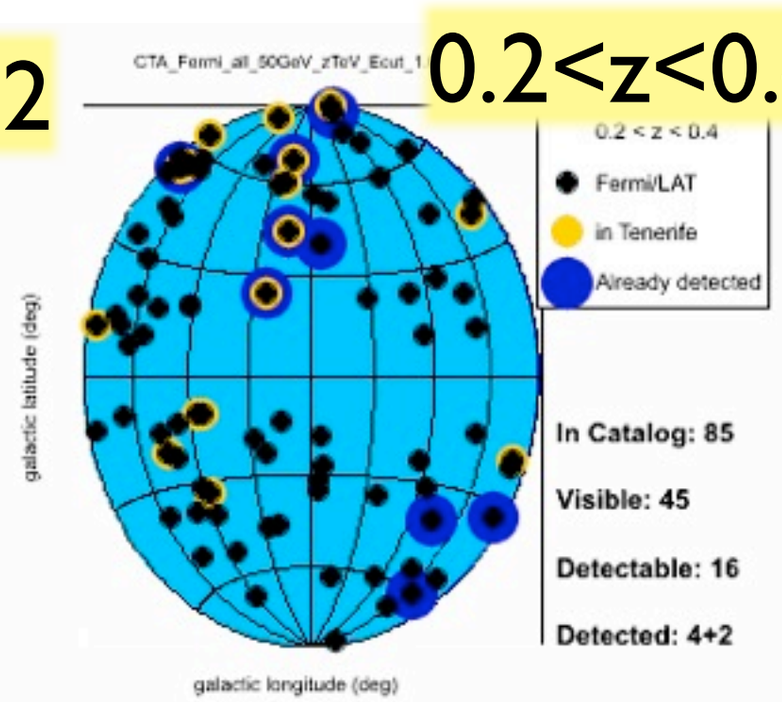


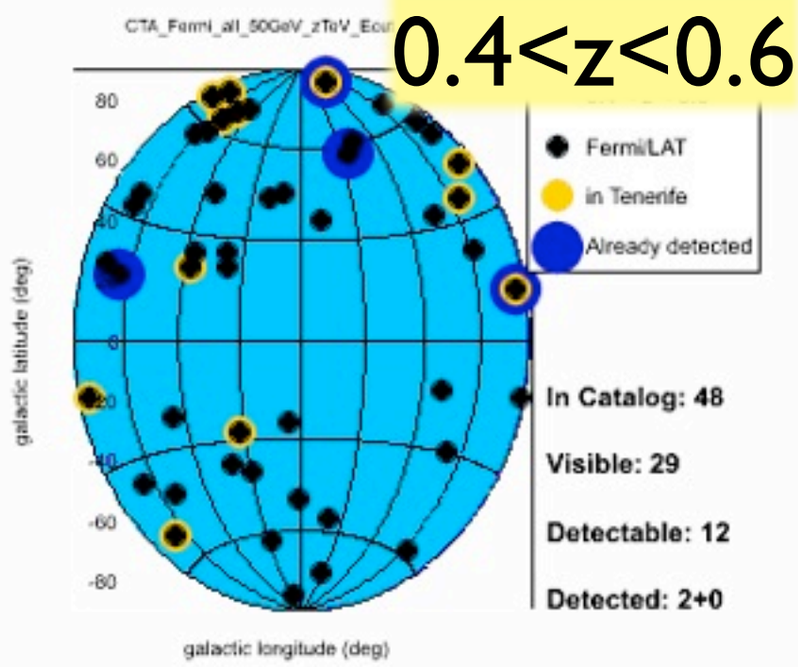
DM & D. Semikoz 2013, in preparation

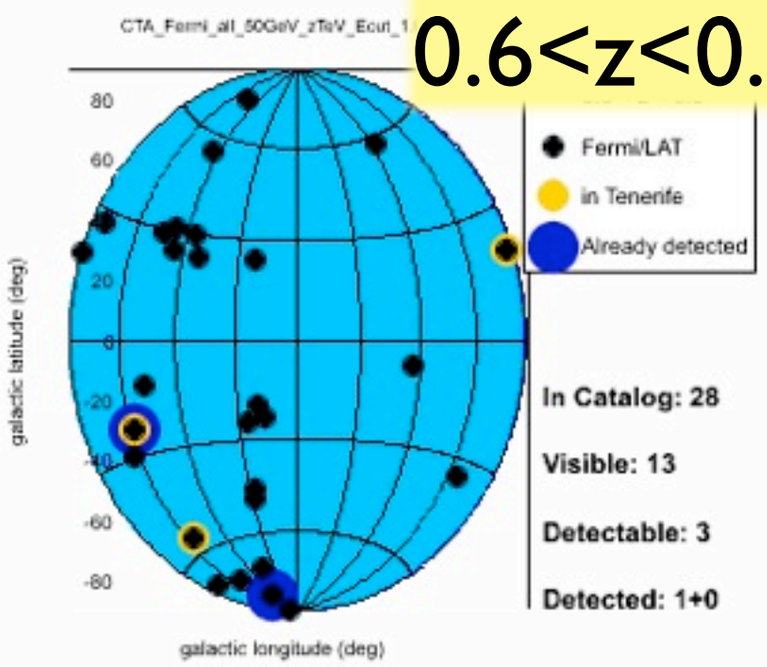
Expectation: CTA in North (20h)

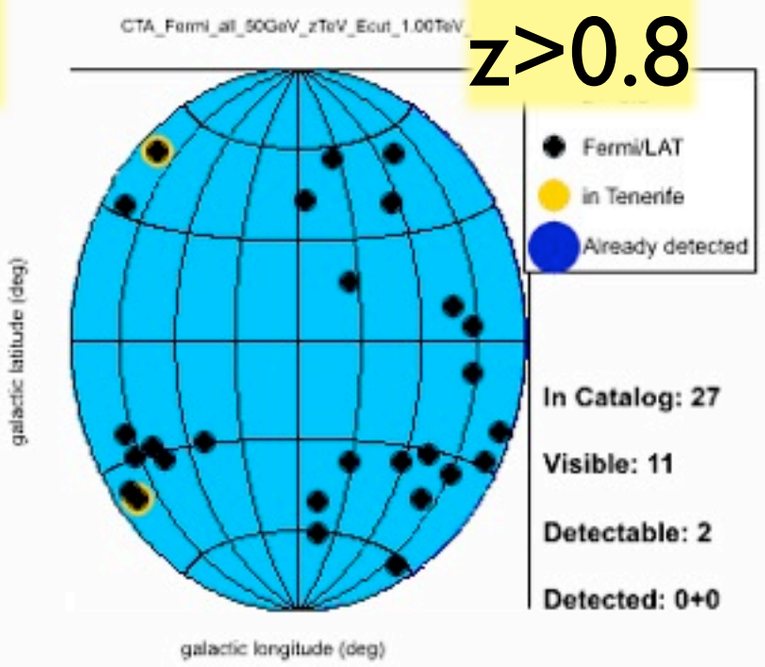












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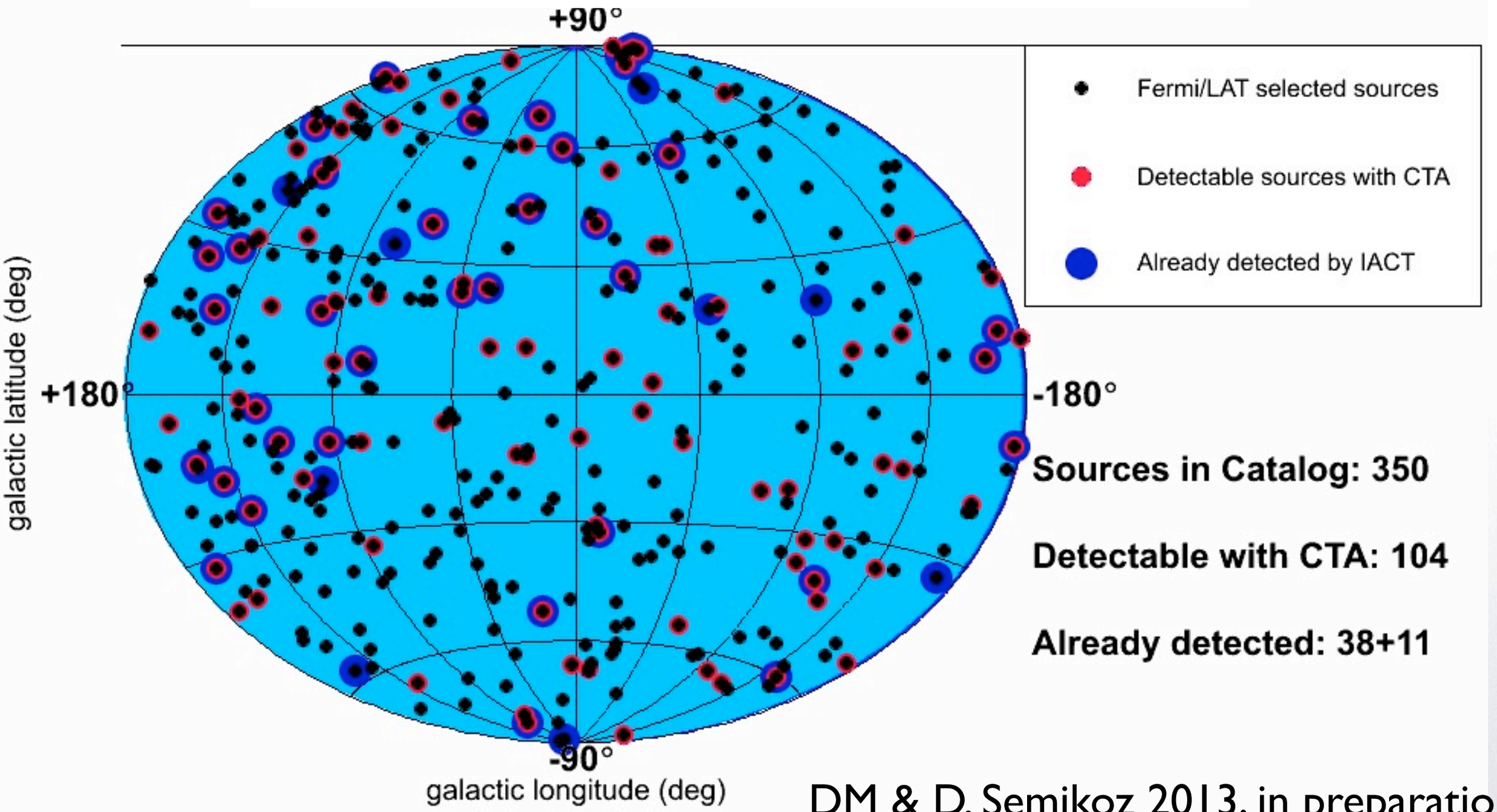
What does it mean for HESS-II?

MAGIC-like everywhere



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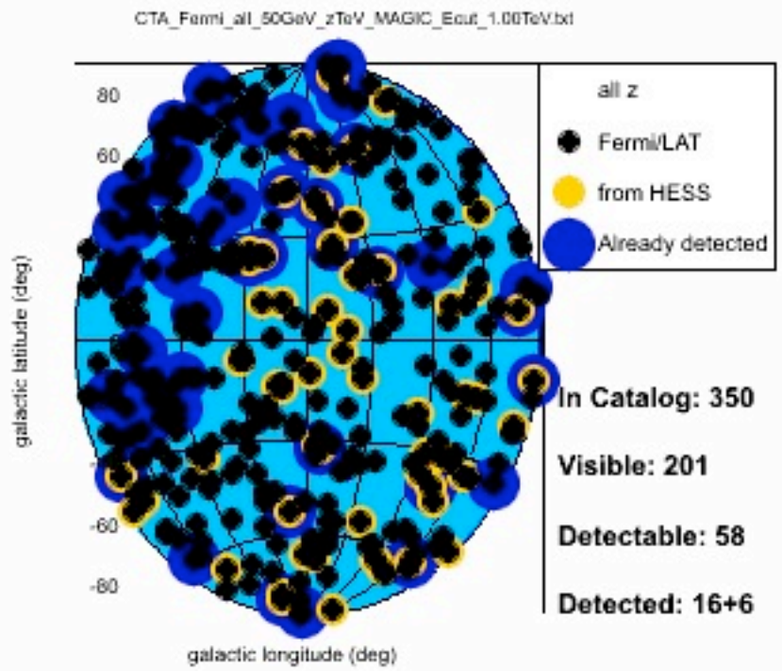
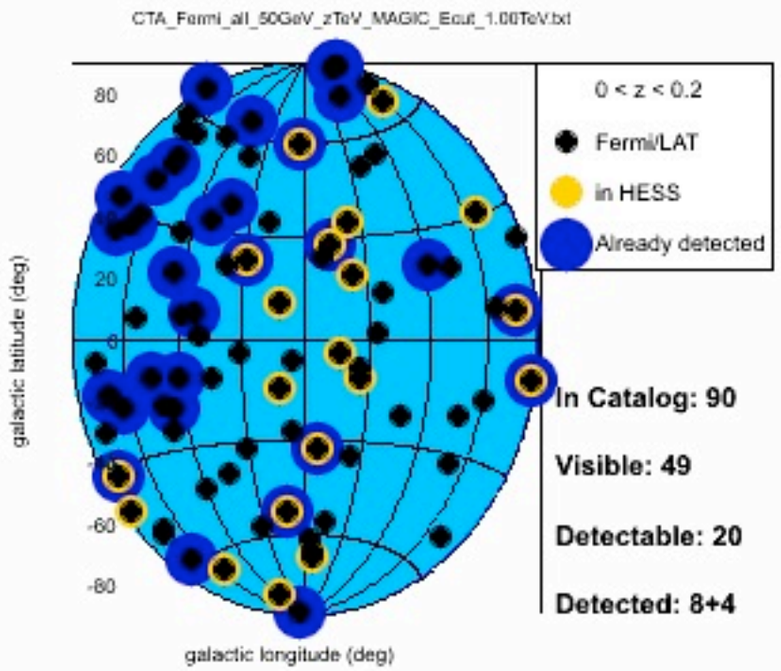
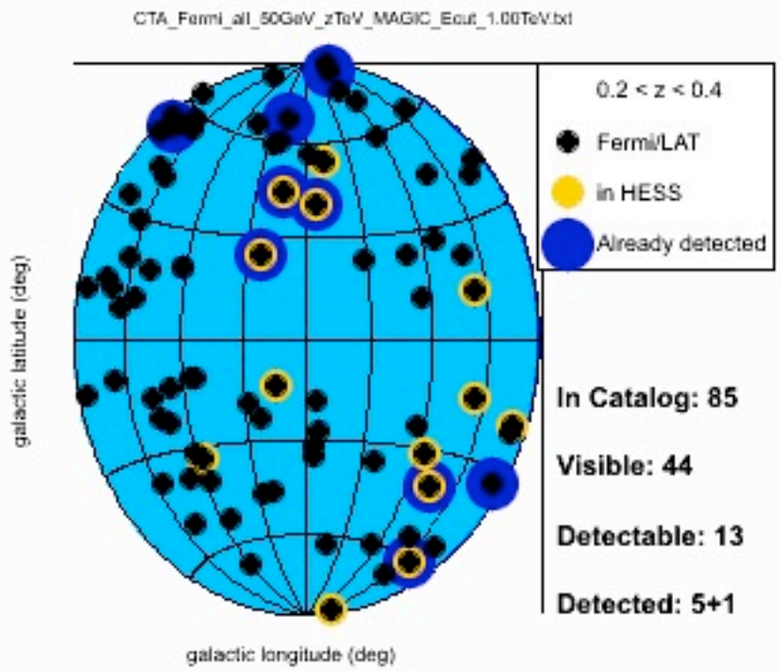
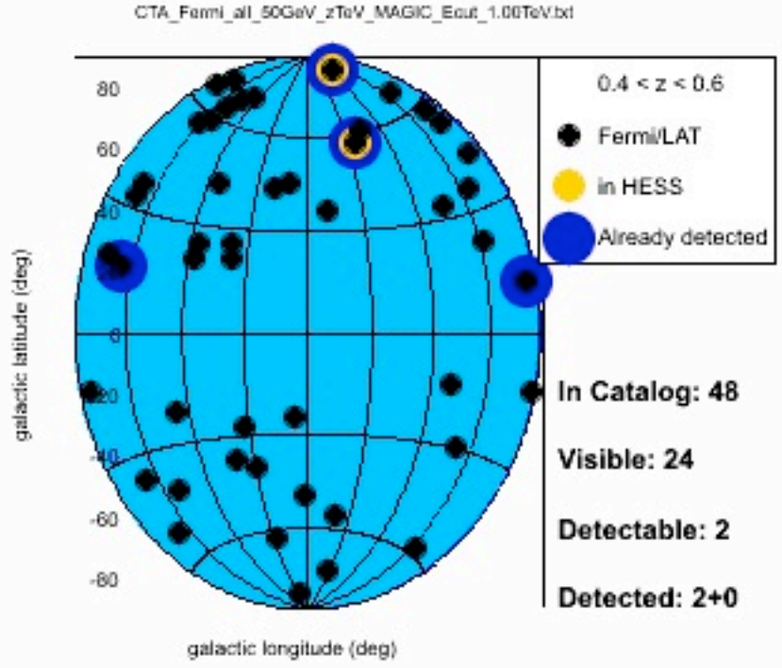
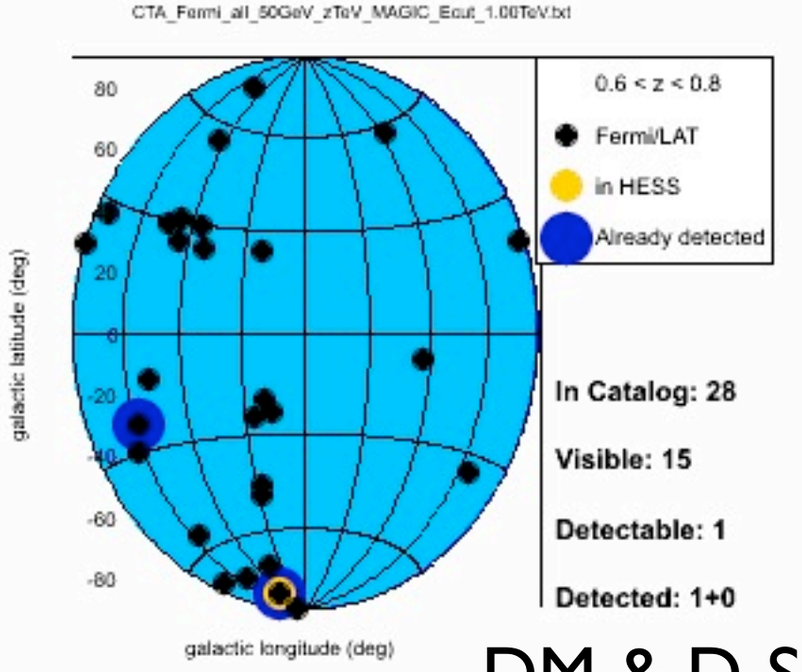
What does it mean for HESS-II?

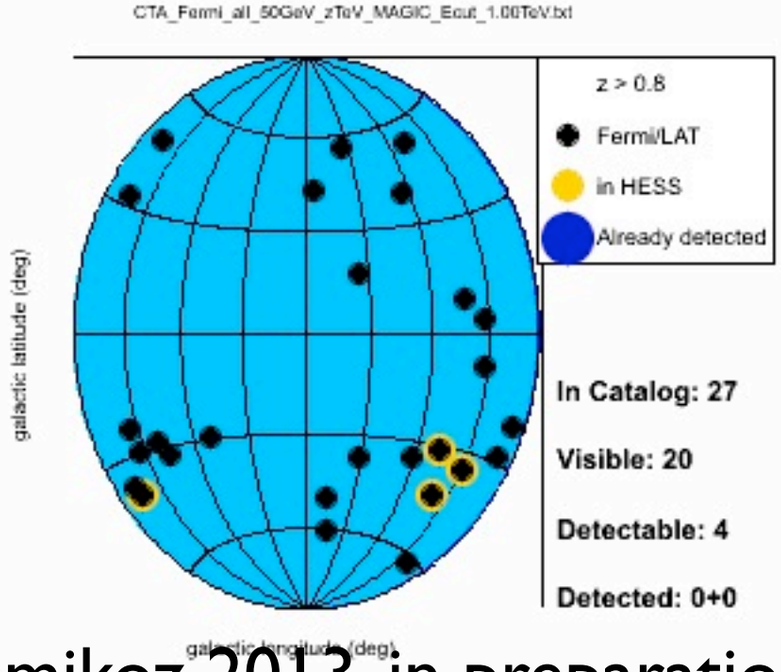












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Best candidates for HESS-II

- PKS 0426-380, z=1.111, E>50GeV = 7
- 1RXS J020922.2-52292 z=0.31, E>50GeV =7
- PKS 0537-441, z=0.892, E>50GeV = 4
- 1RXS J054357.3-55320, z=0.27, E>50GeV = 5
- PMN J1603-4904, z=unknown, E>50GeV = 9

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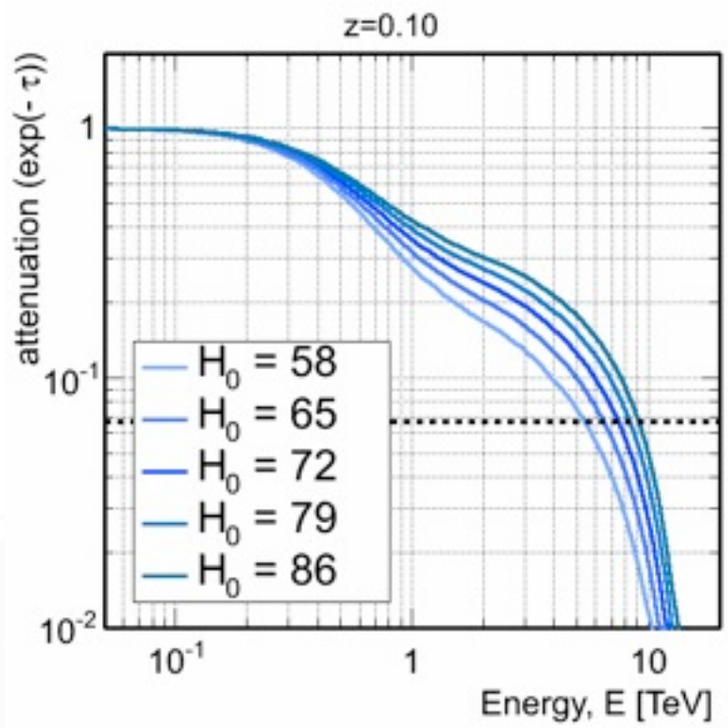
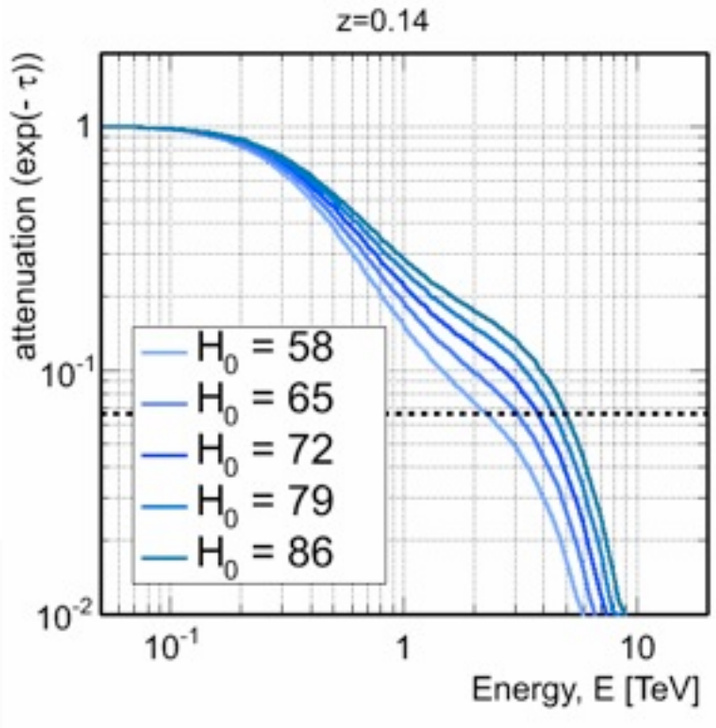
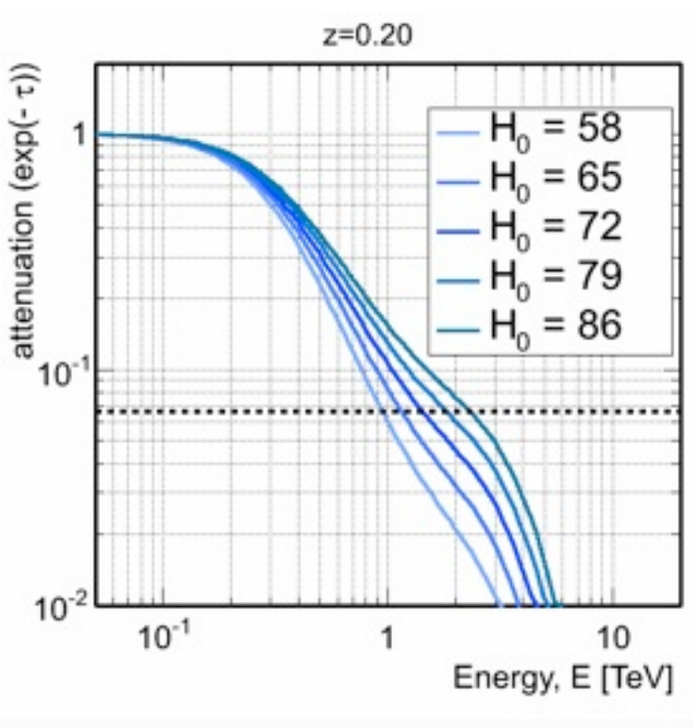
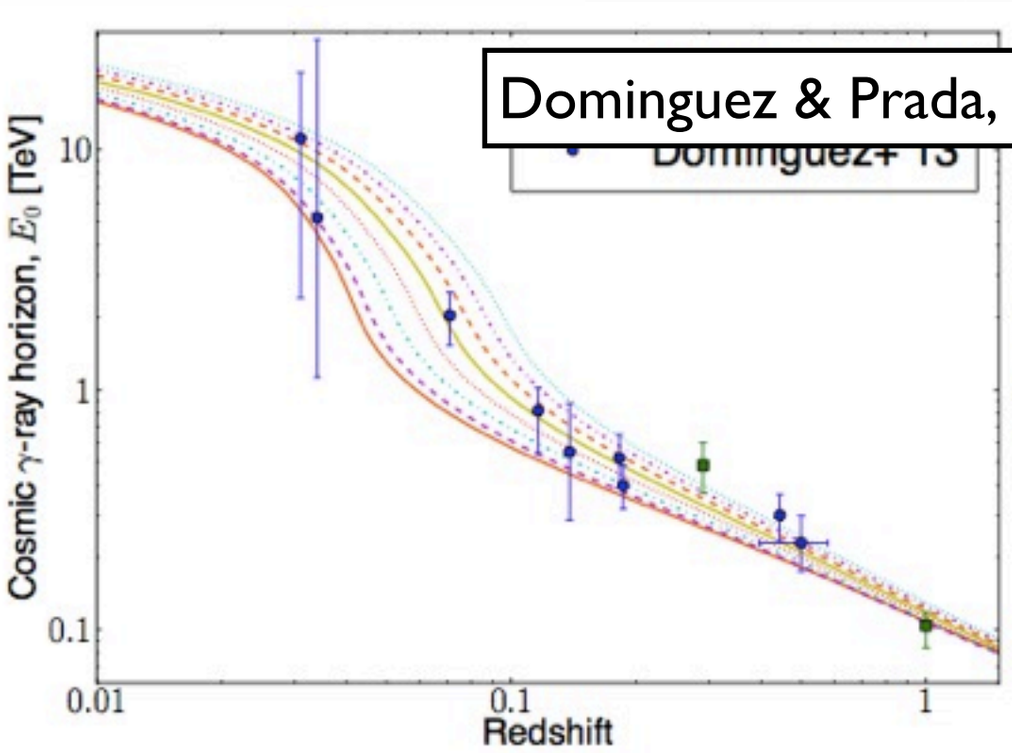
$\tau \sim 1/H_0$ Hubble constant issue

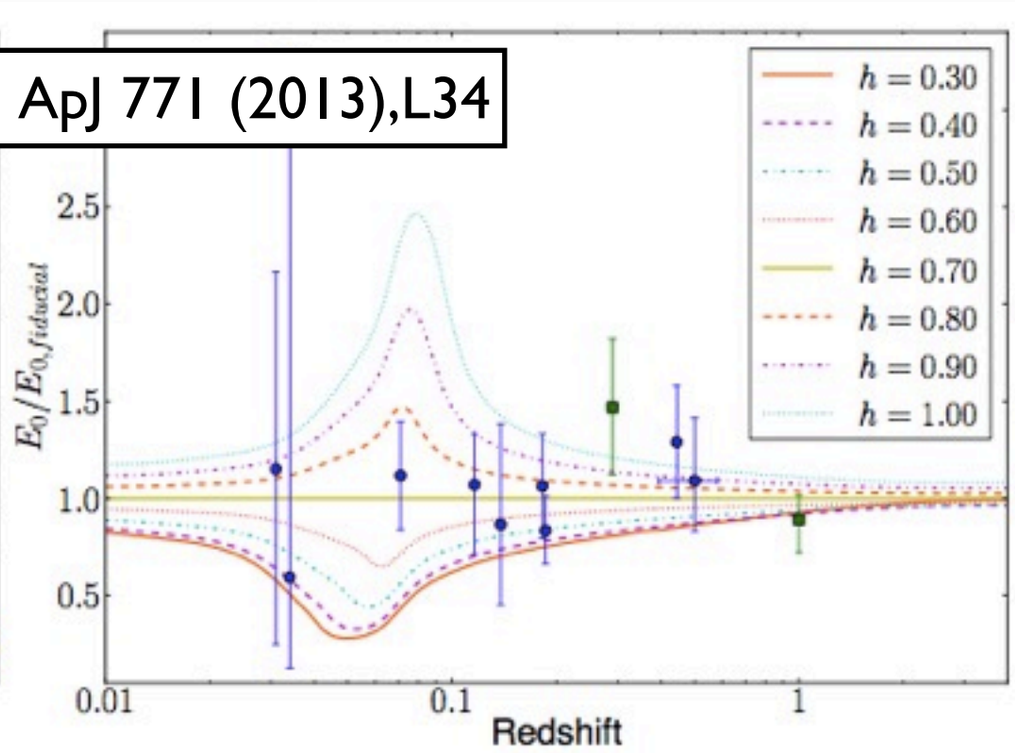






Dominguez & Prada, ApJ 771 (2013),L34





Conclusions

- Good times for EBL constraints
- Even with lowest EBL possible, pile-ups at high energy. This is intriguing / disturbing
- GRBs: CTA/HESS-II can detect them up to z=6 if they are bright enough
- Sources for EBL/cosmology studies:
  - need redshift determination!
  - need monitoring of flaring activity, not many sources that can be detected otherwise
  - some easy catches for HESS-II are there even in steady state
  - good potential to get 1-2 sources beyond z=0.8