Experimental Radiative Near-infrared Atomic Data for Kilonova Spectroscopy

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The Radiative Transfer and Atomic Physics of Kilonovae, Stockholm September 4-7, 2023

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We are focusing on the **measurement** of **TRANSITION PROBABILITIES** (oscillator strengths) of rare earth elements, in particular **Nd III** (Pratyush's PhD) (collaboration with Imperial College London)

OUR CAPABILITIES

- Neutral, singly and doubly ionised atomic spectra (hollow-cathode lamp)
- Spectral range: UV-visible (200 800 nm)
- 1.5 m diffraction grating monochromator (Czerny-Turner, 2400 lines/mm)
- Resolving power: 150 000 (at 450 nm)
- Setting-up a Fabry-Pérot interferometer



Collaborators

Laboratory spectroscopy (Malmö and Lund universities):

Hampus Nilsson, Madeleine Burheim, Asli Pehlivan Rhodin (until 2018), Lars Engström

Calculations (GRASP2k and ATSP2k):

Per Jönsson, Jörgen Ekman, Yanting Li (Malmö University) Jon Grumer, Sema Caliskan (Uppsala University)

Collaborations:

LUMCAS: Wenxian Li, Rickard du Rietz, Stefan Gustafsson, Nils Ryde, Brian Thorsbro, Si Ran, Tomas Brage CompAS; Imperial College London; Atomic Spectroscopy group at NIST Dr Belmonte at Valladolid university, Spain







The Basics

The so-called allowed atomic transitions are between odd and even parity states, and the wavelength is determined by the energy difference.





Close levels - infrared transitions Well separated levels - optical or ultraviolet transitions

Energy level diagram

The energy level diagram can include the physics to understand the spectrum and the spectral lines. $\overline{f_{E}}$





Different needs of radiative data

The need of atomic data in terms of parameters and quality is dependent on the application:

Stellar abundances - individual lines: wavelength, accurate oscillator strengths, line structure (hyperfine structure)

Exoplanet atmosphere detection - wavelength, approximate line strength

Kilonova light curves - broad opacity, completeness

Kilonova spectra - rather accurate wavelength, rather accurate strengths

The near-infrared region (1-5 microns) is particularly empty of atomic data. This was the case for the ultraviolet region before the *Hubble Space Telescope* launch in 1990s.

near-IR atomic transitions

Near-IR transitions (1-5 microns) from an atomic physics point of view, appear differently in different ions and depend on the atomic structure: (remember, near-IR = levels close in energy):

- :: High-excitation transitions; Rydberg states
- :: Lower excitation transitions in some complex atomic spectra
- :: Resonance lines in heavier elements, e.g. third spectra of rare-earth elements



*IUPAC conventional atomic weights; standard atomic weights for these elements are expressed in intervals; see iupac org for an explanation and values.

For a description of the data, visit physics.nist.gov/data NIST SP 966 (September 2014)



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Magnesium



Burheim et al. (2023)

 $10\,000 \text{ cm}^{-1} = 1 \text{ micron}$



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High-excitation states in Fe-group





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ScI, YI and LaI



The energy level structure of Sc I, Y I and La I are similar with the energy levels for nd(n+1)s² and nd²(n+1)s

Sc I : 3d4s², 3d²4s Y I : 4d5s², 4d²5s La I : 5d6s², 5d²6s

Strong lines in infrared of Sc I



Spectra of Galactic center M stars, Showing strong Sc I and Y I lines, but Likely due to low excitation rather than overabundance



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Ce II



The open f-shell gives rise to the many levels and transitions. Resonance transitions in near-IR.

Lanthanides were earlier studied for applications to lighting applications, such as fluorescent tubes.

Lawler et al. (2001)

Lanthanum



For many rare-earth elements, the near-IR transitions include the lowest excitation lines

Lowest levels of Ce III



Term analysis of Zr II



Burheim and Nilsson et al., in progress

Experimental transition rates (or f-values)

The transition rates A (or log gf-values) are derived from two *independently* measured quantities:

- a) The relative emission line intensities from the lines from an common level
- b) The radiative lifetime

$$A_{ul} = \frac{BF_{ul}}{\tau_u}$$

$$BF_{ul} = \frac{I_{ul}}{\sum_{l} I_{ul}}$$



Experimental transition rates (or f-values)

The transition rates A (or log gf-values) are derived from two <i>independently</i>	Decay region	Line spread	Lifetime
 measured quantities: a) The relative emission line intensities from the lines from an common level b) The radiative lifetime 	Only optical	Narrow :)	1-20 ns :)
	Optical+nIR	Broad :(1-20 ns :)
$A_{ul} = \frac{BF_{ul}}{\tau_u}$	Only nIR	Narrow :)	1-10 us :(
$BF_{ul} = \frac{I_{ul}}{\sum_{l} I_{ul}}$			

Lifetime measurements



Laser induced plasma probed with short pulses (1ns)



Measurements of radiative lifetimes

University of Wisconsin (US)

Measurements using one step excitation on ion beam

La I (2015), Sm I (2013), Nd I (2011), Gd I (2011), Er I (2010), Ce I (2009), Er II (2008), Nd II (2003), Eu I (2002), Tb II (2001)

Lund VUV laser lab (SE)

Laser induced fluorescence on laser produced plasma. One or two steps, can reach highly excited levels.

* Jilin University, Changchun (CH)

Experimental setup from the Lund VUV laser lab

Line Intensity measurements

Light source Hollow cathode discharge lamp



Plasma with neutral and singly ionized atoms

Detector Fourier transform spectrometer



Near-IR and optical wavelength Resolving power R=10⁶ Wavelength calibration about 1:10⁷

Emission spectrum of a Cerium discharge



Measurements together with S.Caliskan, J.Grumer, M.Burheim, H.Nilsson

Radiative transfer effect in the discharge - self absorption



Combined approach for larger data sets

Approach for more complete sets of radiative data: Mg I, Al I, Si I/II:

FTS measurements combined with lifetime data

ATSP2k or GRASP2k calculations for additional states and lifetimes

Stellar spectra for benchmarking, application and priorities Linelists such as Gaia-ESO

FTS spectrum of Aluminium



Burheim et al. A&A 672, A197, 2023)

Results for Magnesium



Mg I (Pehlivan Rhodin, et al., 2017)

Results Silicon



C. Sneden, private communication



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NIST SP 966 (September 2014)

Results for Sc I



Pehlivan et al., (2015)

Uncertainties



Experimental f-values can be obtained with uncertainties down to 5%

Uncertainties for theoretical data are harder to estimate, and vary among lines in a calculation.

Parity forbidden lines in near-IR

An important class of infrared lines are parity forbidden transitions (E2 and M1), observed in nebula and low density plasmas.

Low transition rates (A around 1 s⁻¹)

Long radiative lifetimes (several seconds)

Sensitive to collisions

$$A_{ul} = \frac{BF_{ul}}{\tau_u}$$

Can be measured using selective methods at storage rings (e.g. DESIREE @ Stockholm university) combined with astronomical observations of low-density plasmas (Eta Carinae).



Preliminary results for Ba II : $5d^2D_{3/2}$



The pump-probe technique has successfully been developed for the Ba+ ions at DESIREE.

- Ion beam lifetime measured to 500s.
- Effect from repopulation and cascades is very small.
- An uncertainty of a few percent can be reached for ideal systems.
- For $5d^2D_{3/2}$ we reach a lifetime $\tau = 80 \pm 1$ s

Technique will be applied to more complex systems such as FeII and Ni II with astrophysical importance .



Experiments can provide accurate wavelengths for most lines for an ion, also for complex spectra. Level energies require analysis.

Accurate line strengths are derived for selected sets of transitions, to an accuracy of 5% The majority of line intensities are to be provided by calculations. Different approach depending on target ion.

Near-infrared experimental data has been provided for Mg I, Sc I, Al I and Si I and Si II

Analysis in progress: Y I and La I (IR), Zr II

Ongoing: Ce I and Ce II. We investigate an approach to apply a tuning technique to merge the theoretical and experimental linelists, with the extension to Ce III using additional light source.

Forbidden lines can be measured using a stored ion beam.



Term analysis - new levels



Int.	WL/Å	$\rm WN/cm^{-1}$	Lower level	$\rm Energy/cm^{-1}$
32	3654.8387	27353.196	$4d^2(^{3}F)5p z^4D_{3/2}$	59609.941
74	3618.4036	27628.618	$4d^2(^{3}F)5p z^4D_{1/2}$	59609.939
45	3440.3536	29058.452	$4d^2(^3F)5p \ z^4F_{5/2}$	59609.924
82	3426.6525	29174.634	$4d^2(^{3}F)5p \ z^2D_{3/2}$	59609.941
29	3351.1069	29832.311	$4d^2({}^3F)5p z^4F_{3/2}$	59609.939
45	3320.7481	30105.034	$4d^2(^3F)5p \ z^2F_{5/2}$	59609.937
100	3161.0331	31626.072	$4d^2(^{3}F)5p \ z^4G_{5/2}$	59609.940

Hyperfine structure



The effect from hyperfine structure can affect the derived abundance and is more prominent in the near-infrared wavelength region

Laser induced plasma probed with short pulses (1ns)



Advice when asking for atomic data

To maximize the outcome of a request for atomic data:

- Be specific in terms of element, ionization, wavelength region, excitation
- If possible, specify a desired accuracy
- Provide a scientific motivation

No : We need all iron group elements, for spectra I, II and III Yes : We need the oscillator strength for the Ni I line blending the O I 6300Å line to an accuracy of 0.02 dex. It is important for the longstanding controversy of the solar abundance of oxygen.