



# **Preliminary results of Small-scale localized structures in sub-alfvenic regions of solar wind**

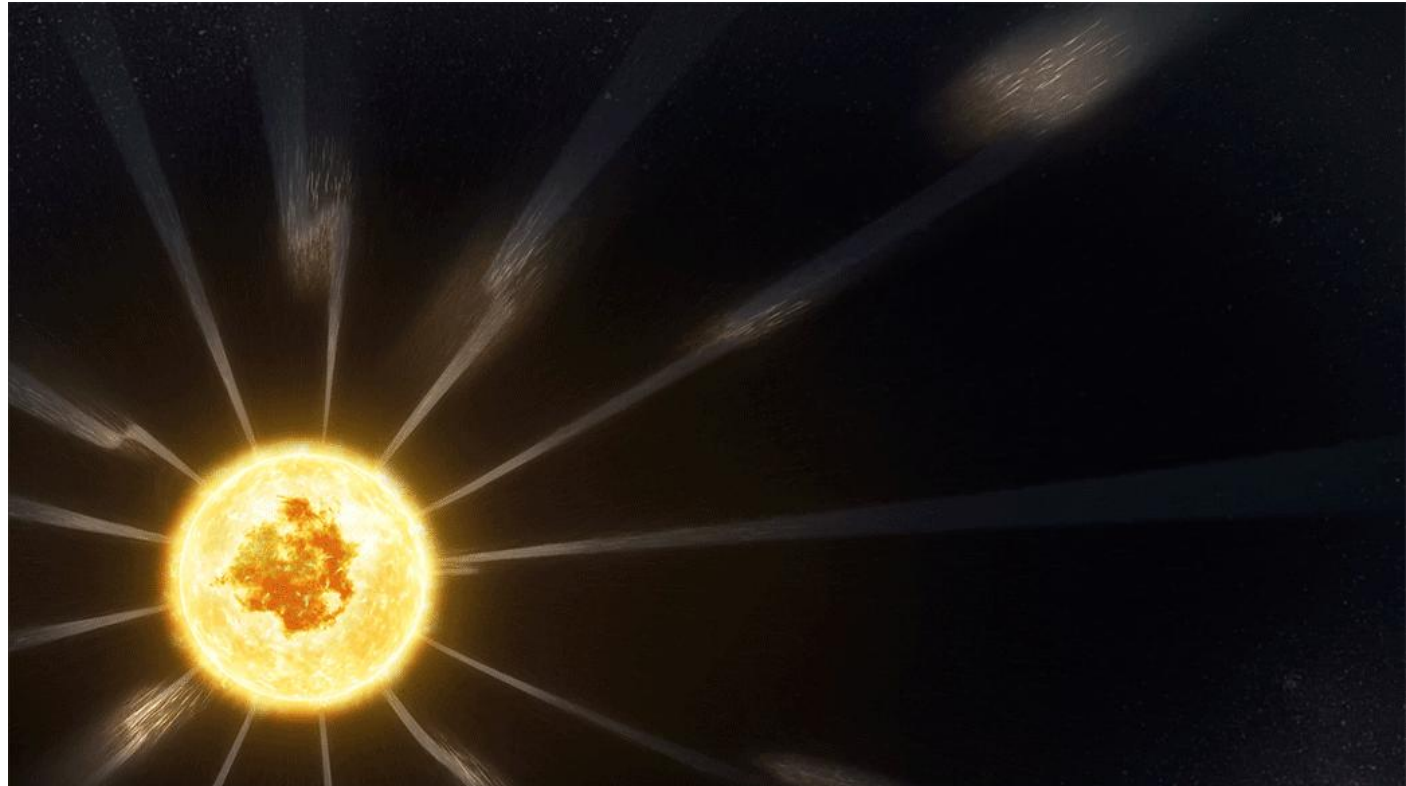
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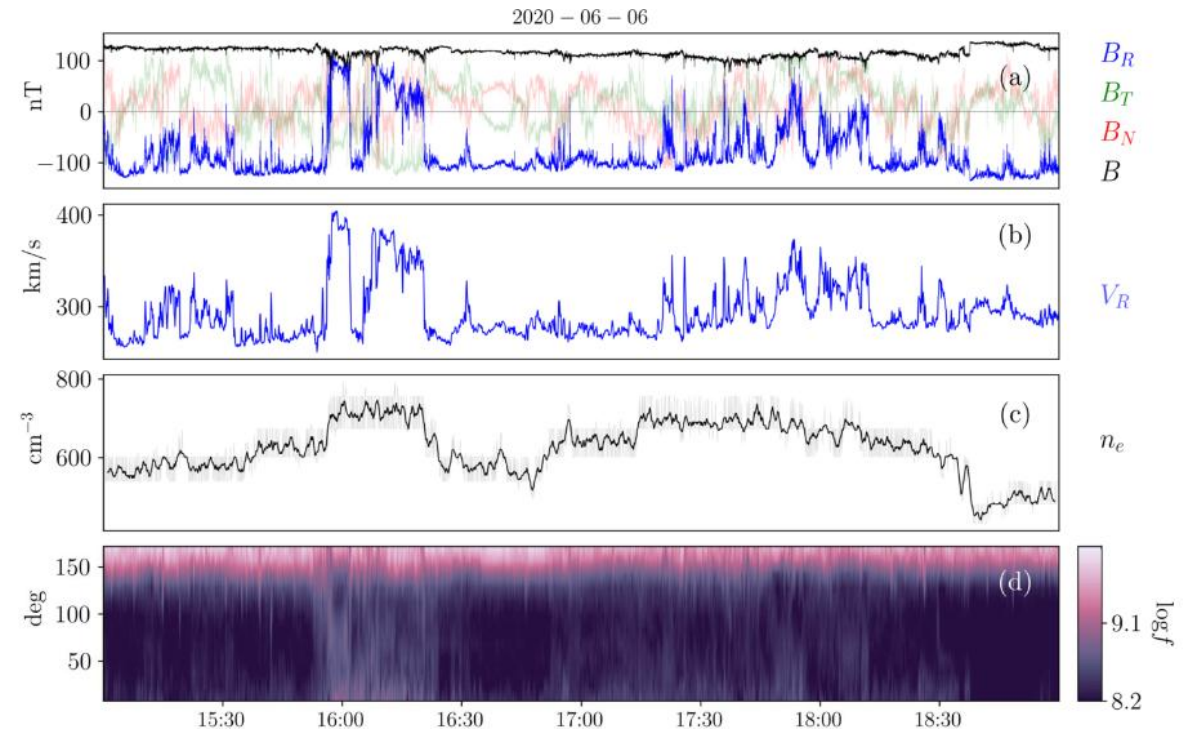
# What are the Switchbacks!

- Large-amplitude, localized deflection of the radial component of the interplanetary magnetic field, that folds back on itself without crossing the heliospheric current sheet.



# What are the Switchbacks!

- Rapid reversal of the radial magnetic field (often  $>90^\circ$  rotation).
- Velocity enhancement (jets).
- No change in the Pitch Angle Distribution PAD.
- Alfvénic.

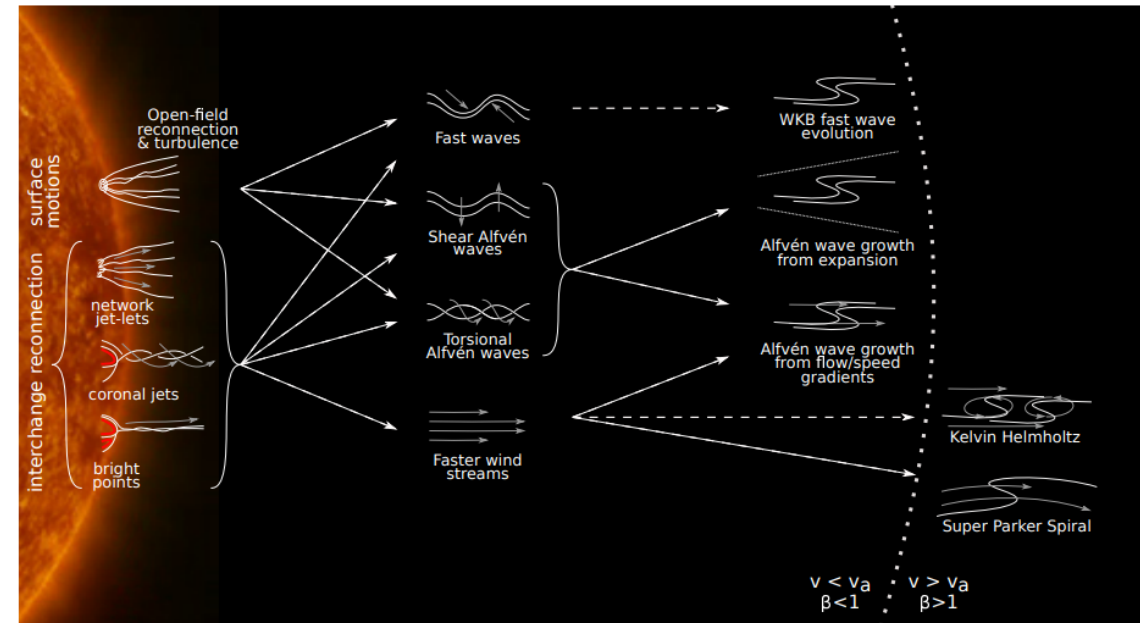


# History!

- They were first detected by Ulysses as 'polarity inversions' (Balogh et al. 1999).
- large-scale magnetic field inversions at sector boundaries observation using WIND (Crooker et al. 2004).
- Using Ulysses data, first used terminology 'switchbacks', Yamauchi et al. (2004).

# How they are formed!

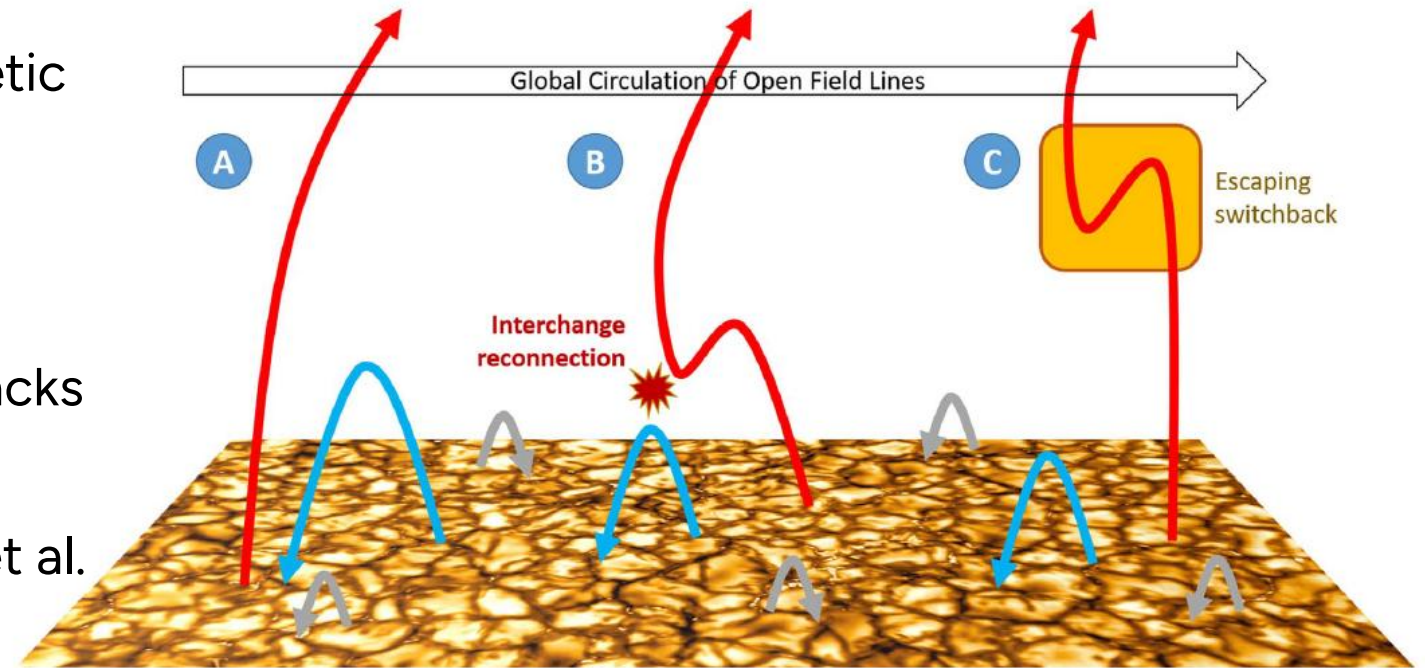
- **In Lower Corona:** Switchbacks originate near the Sun through magnetic reconnection or photospheric processes; they propagate outward as pre-formed structures.
- **In Situ:** Switchbacks are generated locally in the expanding solar wind through plasma instabilities, turbulence, or expansion effects
- Neither paradigm alone is sufficient.



# Interchange Reconnection

- Open magnetic field lines reconnect with closed coronal loops near the solar surface
- This process launches kinked magnetic field lines and plasma jets into the heliosphere
- Interchange reconnection can eject helical flux ropes into the solar wind, which may be identified as switchbacks

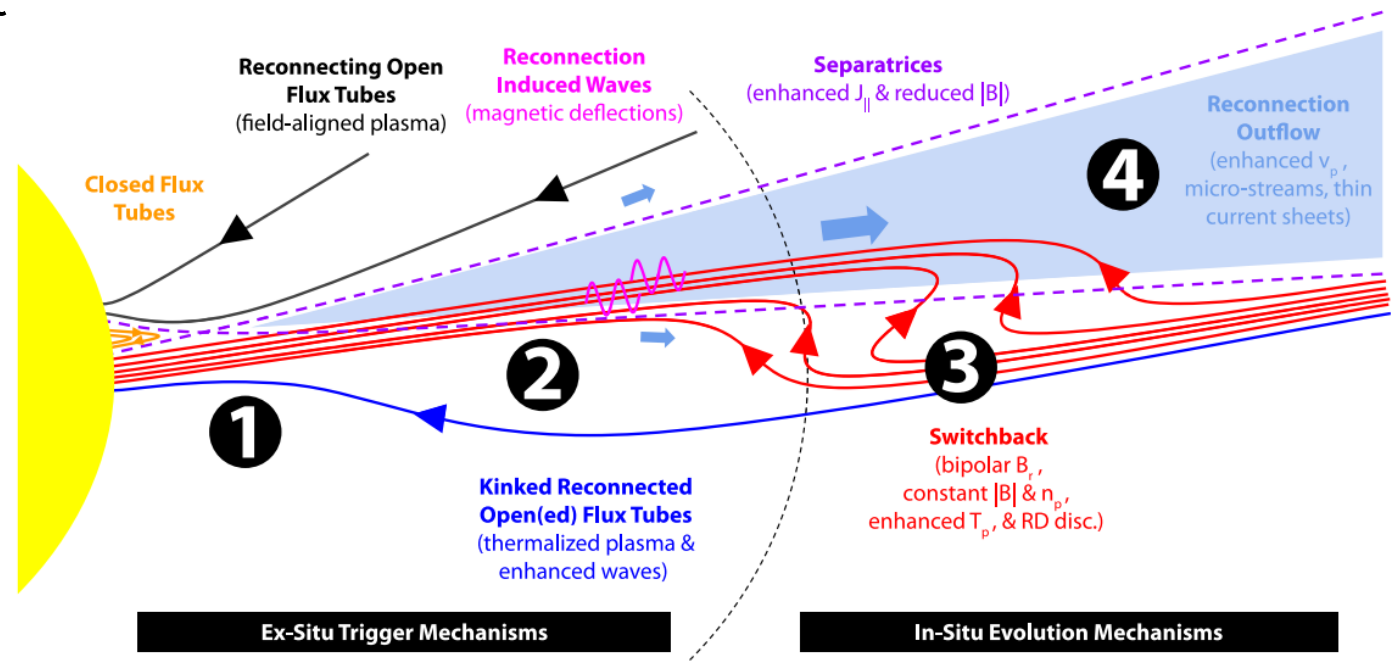
(Fisk 2005, Fisk & kasper 2020, Zank et al. 2020, Bale et al. 2021)



Fisk & Kasper 2020

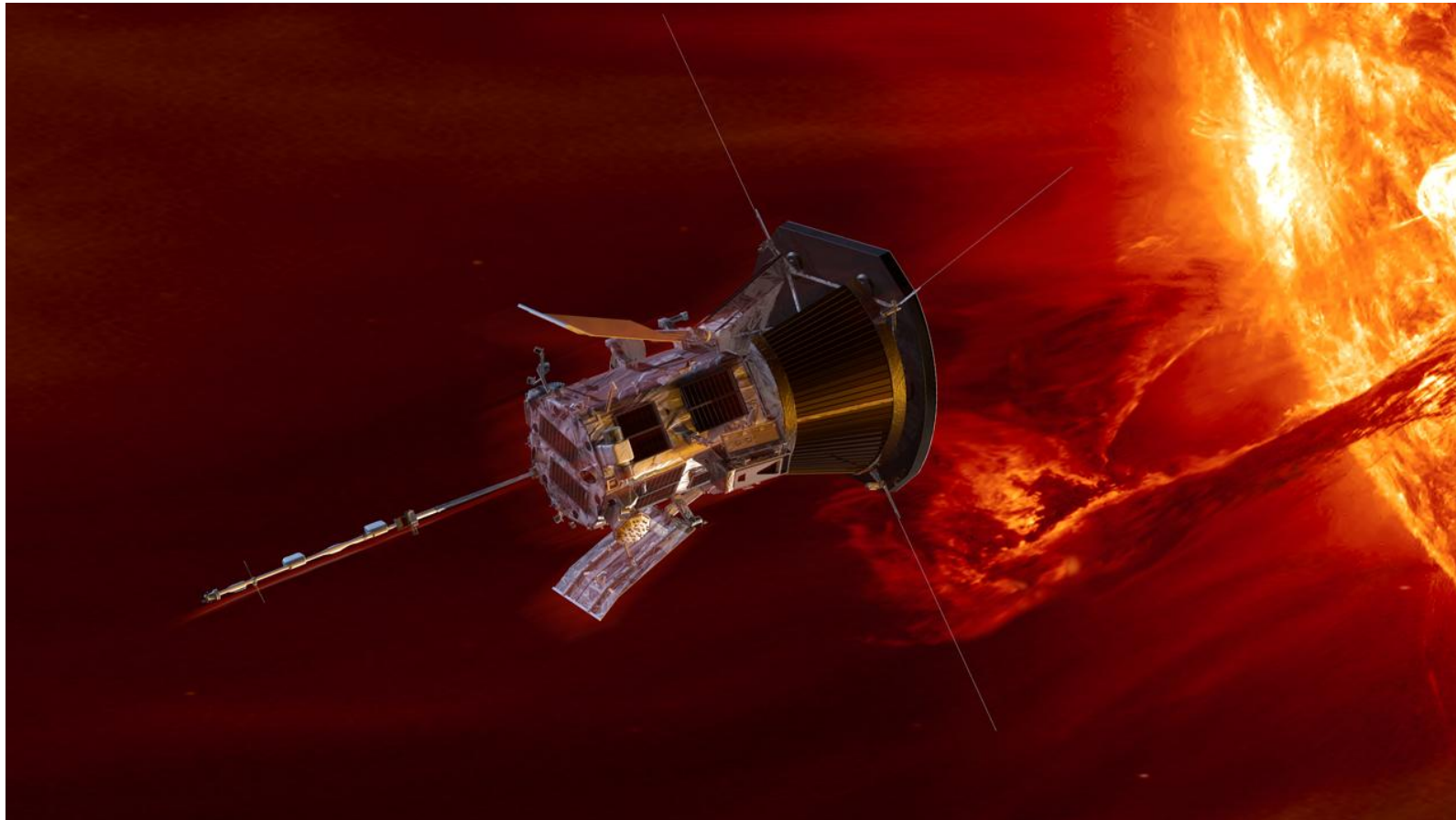
# In situ mechanisms

- Alfvénic turbulence in the expanding solar wind can spontaneously generate magnetic reversals.
- Flow shear across reconnection outflow boundaries can trigger instabilities that grow magnetic deflections
- Kelvin-Helmholtz instability at switchback boundaries may play a role in their evolution



# PSP... Finally!

- observations within 0.3 AU showed:
- Switchbacks occur in both slow and fast solar wind streams
- prevalence decreases with heliocentric distance
- no characteristic magnitude or duration



# PSP ... Unfortunately!

- The rate of occurrence falls off sharply approaching the Sun near 0.2 au (40 Re) and rises gently from 0.2 au outward.

(Pecora et al. 2022)

- enhanced magnetic fluctuations ( $\theta_B \leq 90^\circ$ ).
- A clear absence of radial field reversals ( $\theta_B \geq 90^\circ$ ) in the sub-Alfvénic corona.

(M. Akhavan-Tafti and S. L. Soni, 2024)

**Table 1**  
List of PSP  $R_A$  Crossings and  $B_r$  Reversal Statistics

Encounter #	$R_A$ crossings (count)	Duration (s) (min)	$B_r$ reversals (count)
1	0	0	663
2	0	0	413
3	0	0	...
4	0	0	449
5	0	0	420
6	0	0	512
7	0	0	286
8	2	140, 90	370
9	1	176	305
10	...	...	329
11	0	0	295
12	3	100, 50, 40	199
13	4	35, 40, 250, 15	246
14	0	0	314

M. Akhavan-Tafti and S. L. Soni, 2024

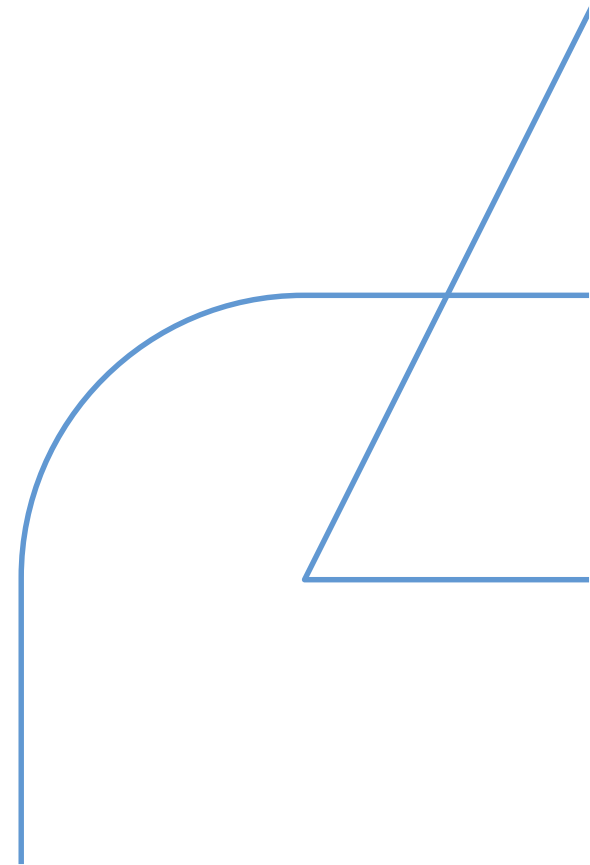
More than 4800  $B_r$  reversal events were identified, about 10% of which were further characterized as switchbacks.

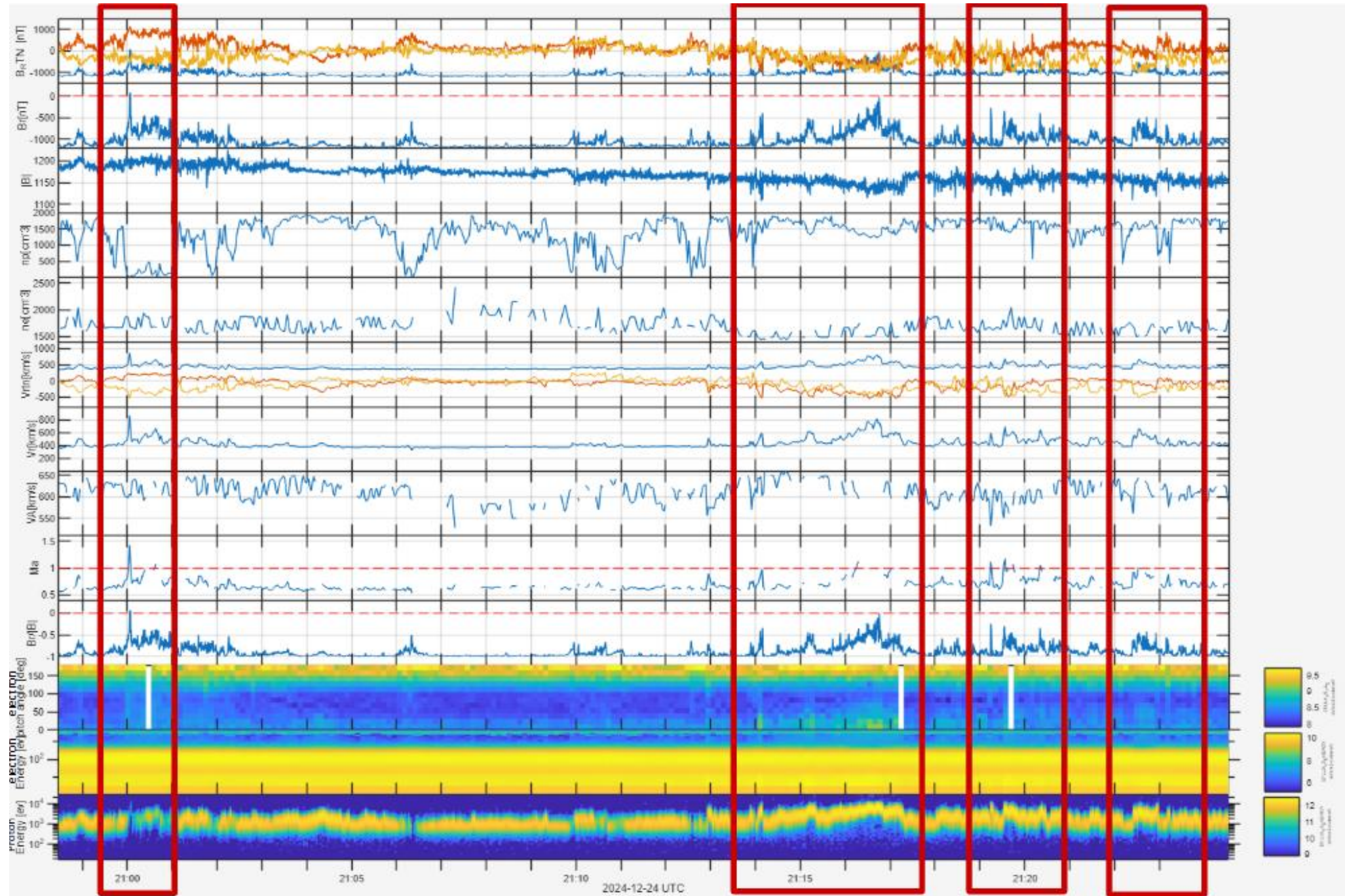
# Why Switchbacks!

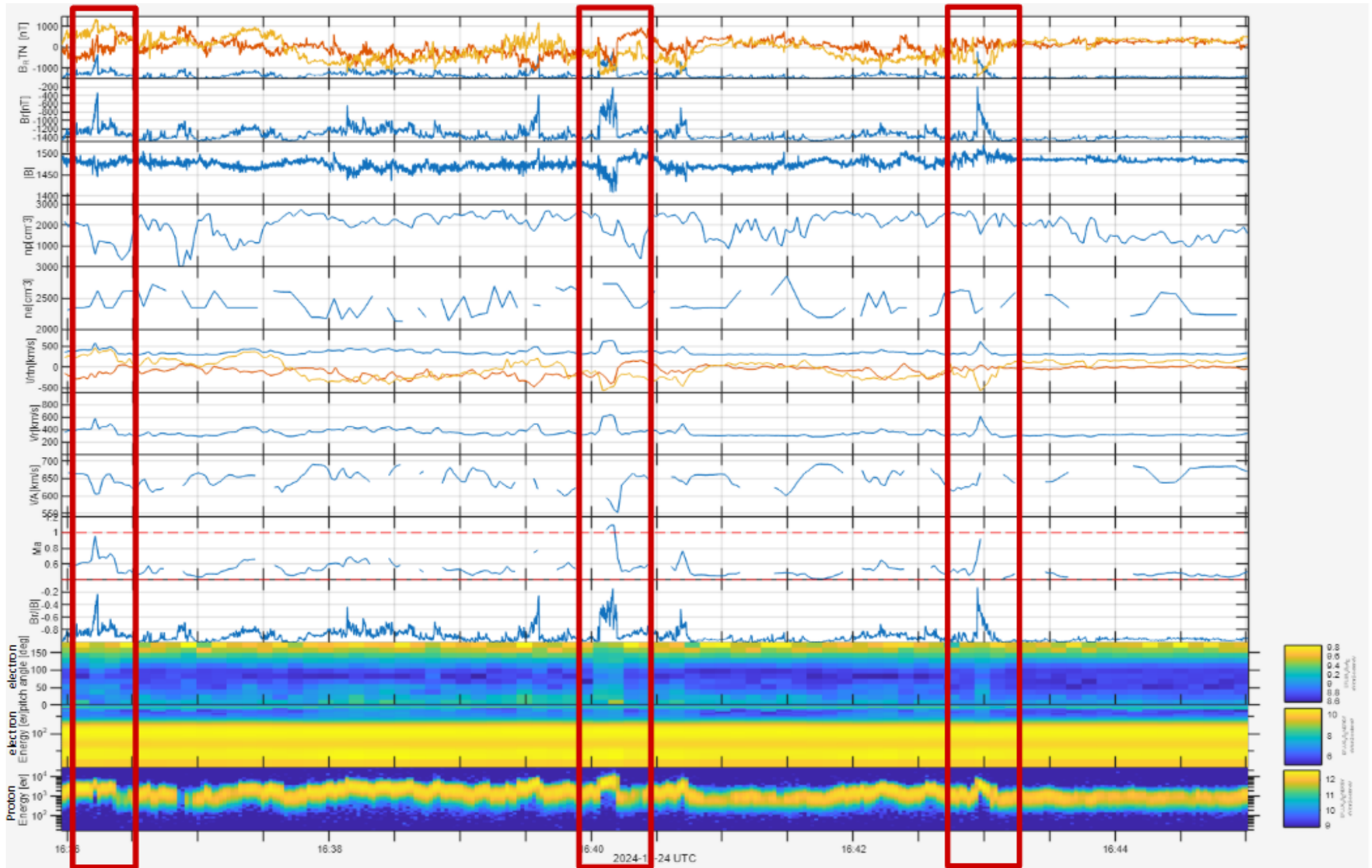
- Unsolved structure with a lot of open questions
  - Where do they form? (Solar surface vs. in situ)
  - How do they evolve with radial distance?
  - Do they contribute to solar wind acceleration?

**Synergies!**

# Preliminary Result from PSP E22









**Thank you!**

